

# The SLUGGS survey: globular cluster stellar population trends from weak absorption lines in stacked spectra

Christopher Usher,<sup>1\*</sup> Duncan A. Forbes,<sup>1</sup> Jean P. Brodie,<sup>2</sup> Aaron J. Romanowsky,<sup>2,3</sup> Jay Strader,<sup>4</sup> Charlie Conroy,<sup>2</sup> Caroline Foster,<sup>5</sup> Nicola Pastorello,<sup>1</sup> Vincenzo Pota<sup>1</sup> and Jacob A. Arnold<sup>2</sup>

<sup>1</sup>Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Hawthorn, VIC 3122, Australia

<sup>2</sup>University of California Observatories, 1156 High Street, Santa Cruz, CA 95064, USA

<sup>3</sup>Department of Physics and Astronomy, San José State University, One Washington Square, San Jose, CA 95192, USA

<sup>4</sup>Department of Physics and Astronomy, Michigan State University, East Lansing, MI 48824, USA

<sup>5</sup>Australian Astronomical Observatory, PO Box 915, North Ryde, NSW 1670, Australia

Accepted 2014 September 30. Received 2014 September 12; in original form 2014 May 15

## ABSTRACT

As part of the SAGES Legacy Unifying Globulars and GalaxieS (SLUGGS) survey, we stack 1137 Keck DEIMOS (Deep Imaging Multi-Object Spectrograph) spectra of globular clusters from 10 galaxies to study their stellar populations in detail. The stacked spectra have median signal-to-noise ratios of  $\sim 90 \text{ \AA}^{-1}$ . Besides the calcium triplet, we study weaker sodium, magnesium, titanium and iron lines as well as the  $H\alpha$  and higher order Paschen hydrogen lines. In general, the stacked spectra are consistent with old ages and a Milky Way-like initial mass function. However, we see different metal line index strengths at fixed colour and magnitude, and differences in the calcium triplet–colour relation from galaxy to galaxy. We interpret this as strong evidence for variations in the globular cluster colour–metallicity relation between galaxies. Two possible explanations for the colour–metallicity relation variations are that the average ages of globular clusters vary from galaxy to galaxy or that the average abundances of light elements (i.e. He, C, N and O) differ between galaxies. Stacking spectra by magnitude, we see that the colours become redder and metal line indices stronger with brighter magnitudes. These trends are consistent with the previously reported ‘blue tilts’ being mass–metallicity relations.

**Key words:** globular clusters: general – galaxies: abundances – galaxies: star clusters: general – galaxies: stellar content.

## 1 INTRODUCTION

Most studies (e.g. Peng et al. 2006) of globular clusters (GCs) beyond the Local Group use integrated optical colours to determine the GC metallicities. Such studies typically assume that all GCs in all galaxies follow the same colour–metallicity relationship. However, the stellar population properties of early-type galaxies scale with galaxy mass (e.g. Terlevich & Forbes 2002; Thomas et al. 2005; Gallazzi et al. 2006; Kuntschner et al. 2010) with age and metallicity increasing with galaxy mass. The abundances of the  $\alpha$ -elements (except Ca), C, N and Na increase with stellar mass (Johansson, Thomas & Maraston 2012; Conroy, Graves & van Dokkum 2014) while the initial mass function (IMF) becomes more bottom heavy (dominated by dwarf stars) with increasing galaxy mass (e.g. Con-

roy & van Dokkum 2012b; La Barbera et al. 2013) in line with evidence from dynamical studies (e.g. Cappellari et al. 2012).

Since the spatial distribution (e.g. Strader et al. 2011b; Forbes, Ponman & O’Sullivan 2012) and kinematics (e.g. Pota et al. 2013a) of red GCs are observed to agree well with those of their host galaxy’s starlight, the GC stellar population parameters are also expected to scale with galaxy mass. The mean GC colour becomes redder with brighter galaxy luminosity, seeming to indicate that the average GC metallicity increases with galaxy mass (e.g. Brodie & Huchra 1991; Peng et al. 2006; Strader et al. 2006). Since the ages, chemical abundances and IMF of a stellar population affect its integrated colours (e.g. Conroy, Gunn & White 2009), variations in the ages, chemical abundances or IMF of GCs between galaxies would lead to variations in the colour–metallicity relation between galaxies.

There have been only a few studies comparing the GC colour–metallicity relation between galaxies. Vanderbeke et al. (2014b)

\*E-mail: [cusher@astro.swin.edu.au](mailto:cusher@astro.swin.edu.au)

compared the colours and metallicities of Milky Way GCs with GCs from M49 and M87, finding agreement between the three galaxies. However, their comparison lacked Milky Way GCs more iron rich than  $[\text{Fe}/\text{H}] = -0.3$  and had only a couple of M49 and M87 GCs more iron poor than  $[\text{Fe}/\text{H}] = -1.5$ . In contrast, Foster et al. (2011) found a large difference in the GC relationship between colour and the metallicity sensitive Ca II triplet index (CaT) between NGC 1407 and NGC 4494. With a larger data set of 11 galaxies, Usher et al. (2012) saw comparable differences in the GC colour–CaT relationship between galaxies. Assuming that the CaT accurately traces metallicity this would indicate that the colour–metallicity relation varies from galaxy to galaxy. Even if the colour–metallicity relation was universal, a varying CaT–metallicity relation would indicate varying GC stellar populations. Additionally, several studies have found variations in the ultraviolet colour–metallicity relation, with GCs in M87 (Sohn et al. 2006) showing bluer ultraviolet colours than Milky Way GCs, M31 GCs (Rey et al. 2007) showing similar colours to the GCs of the Milky Way and GCs associated with the Sagittarius dwarf galaxy showing redder ultraviolet colours (D’Alessandro et al. 2012).

GCs in most massive galaxies (but not all, i.e. NGC 4472, Strader et al. 2006) show a colour–magnitude relation known as the ‘blue tilt’ where the blue GCs become redder at brighter magnitudes (Harris et al. 2006; Spitler et al. 2006; Strader et al. 2006). The evidence for the red GCs having a colour–magnitude relation is much weaker with Harris (2009) finding no evidence for a ‘red tilt’ while Mieske et al. (2010) found evidence for a weak red tilt only in the centres of high-mass galaxies. Since colour traces metallicity and absolute magnitude traces mass, the blue tilt has usually been explained as a mass–metallicity relationship where the more massive GCs were able to retain more of the metals from the first generation of their stars and self-enrich (Strader & Smith 2008; Bailin & Harris 2009). The population of stars responsible for this self-enrichment should leave a signature in the form of chemical abundance trends with magnitude.

All studied Milky Way GCs show star to star variations in their light element abundances (see review by Gratton, Carretta & Bragaglia 2012). In each GC, significant numbers of stars show enhanced He, N, Na and Al and reduced O, C and Mg abundances compared to the remaining stars which show a ‘normal’  $\alpha$ -element enhanced abundance pattern. There is some evidence (Carretta 2006; Carretta et al. 2010) that the  $[\text{O}/\text{Na}]$  and  $[\text{Mg}/\text{Al}]$  ranges become wider at brighter GC luminosities, while Schiavon et al. (2013) found that metal-rich M31 GCs have higher  $[\text{N}/\text{Fe}]$  at higher masses. These variations in light element are also commonly seen as split or broadened main sequences, subgiant branches or red giant branches (RGBs) in high-quality colour–magnitude diagrams (e.g. Monelli et al. 2013). A few massive MW GCs including  $\omega$  Cen (Freeman & Rodgers 1975; Norris & Da Costa 1995), M22 (Marino et al. 2011a), NGC 1851 (Carretta et al. 2011) and NGC 5824 (Da Costa, Held & Saviane 2014) show significant  $[\text{Fe}/\text{H}]$  dispersions. Several models (e.g. Decressin et al. 2007; D’Ercole et al. 2010; Conroy & Spergel 2011) have been proposed to explain the observed abundance spreads, but a common feature is that GCs were able to self-enrich, with the more massive GCs being able to retain more of their metals. What if any connection exists between the self-enrichment process that creates the blue tilt and the origin of the light element anticorrelations observed within GCs is unclear. However, in  $\omega$  Cen the Na–O anticorrelation is observed over a wide range of iron abundances (Marino et al. 2011b), suggesting that the origin of the blue tilt and the light element anticorrelations do not share a common origin.

Here, we stack large numbers of GC spectra observed with the Deep Imaging Multi-Object Spectrograph (DEIMOS; Faber et al. 2003) on the Keck II telescope to reach sufficient signal-to-noise ratios (S/N) to study their stellar populations in greater detail. A highly multiplexed spectrograph, DEIMOS is most sensitive in the red where the old stellar populations of GCs are brightest. The strongest spectral feature in the DEIMOS wavelength range is the CaT at 8498, 8542 and 8662 Å which has been used extensively to measure the metallicities of both resolved stars and integrated light (see Usher et al. 2012 and references therein). Recently interest in the red wavelength range has focused on using IMF sensitive features such as the Na I doublet at 8183 and 8194 Å and the CaT to study the low-mass IMF in early-type galaxies (e.g. Conroy & van Dokkum 2012b; La Barbera et al. 2013). Other important spectral features include the Mg I line at 8807 Å and several TiO molecular bands (Cenarro et al. 2009). The DEIMOS spectra frequently cover H $\alpha$  as well as the higher order Paschen hydrogen lines. In addition to studying the CaT strengths of GCs, Foster et al. (2010) studied the strengths of the Paschen lines and a handful of weak Fe and Ti in the vicinity of the CaT. Kirby, Guhathakurta & Sneden (2008) and Kirby et al. (2011a, 2011b) have used DEIMOS in a similar setup to ours to measure Fe and  $\alpha$ -element abundances of RGB stars in Local Group dwarf galaxies using weak metal lines.

In Section 2, we describe our sample of galaxies and GCs. In Section 3, we describe our spectral stacking procedure and our measurement of spectral indices. In Section 4, we compare our colour and spectral index measurements with single stellar population models and consider how these spectral indices and colours vary from galaxy to galaxy as well as how they vary with GC luminosity. Finally, in Section 5, we give our conclusions. Throughout this paper, we will use metallicity ( $[Z/\text{H}]$ ) to refer to the total abundances of all metals. This is distinct from the iron abundance ( $[\text{Fe}/\text{H}]$ ), a distinction not always made by all authors. Complicating any comparison between studies is the fact that different models and calibrations are tied to different metallicity or iron abundance scales (for example see Usher et al. 2012).

## 2 DATASET

Here, we study GC spectra from 10 galaxies that are a part of the SAGES Legacy Unifying Globulars and GalaxieS (SLUGGS) survey<sup>1</sup> (Brodie et al. 2014) sample. The SLUGGS survey is an ongoing study of 25 massive, nearby, early-type galaxies and their GC systems using Keck DEIMOS spectroscopy and wide-field imaging from Suprime-Cam (Miyazaki et al. 2002) on the Subaru telescope. Details of the galaxies in this study are presented in Table 1. We used spectra and CaT strengths from Usher et al. (2012) for the eight galaxies in Usher et al. (2012) with more than 40 GCs with CaT index measurements. For these galaxies, we use the radial velocities from Pota et al. (2013a). For NGC 4278, we supplemented this with spectra, CaT strengths and radial velocities from Usher et al. (2013). We add two other galaxies, NGC 4473 (Pota et al. 2013b; Pota et al. in preparation) and NGC 4649 (Pota et al. in preparation), for which we have large numbers of DEIMOS spectra.

The spectra were all observed with DEIMOS in multislit mode between 2006 and 2013 with the primary aim of measuring GC radial velocities. Exposure times averaged two hours per slit mask. All observations used a central wavelength of 7800 Å, the 1200 line  $\text{mm}^{-1}$  grating and 1 arcsec slits. This setup yields a resolution of

<sup>1</sup> <http://sluggs.swin.edu.au>

**Table 1.** Galaxy properties.

Galaxy	$V_{\text{sys}}$ (km s <sup>-1</sup> )	$D$ (Mpc)	$M_K$ (mag)	Spectra	GCs
(1)	(2)	(3)	(4)	(5)	(6)
NGC 1407	1779	26.8	-25.4	223	203
NGC 2768	1353	21.8	-24.7	52	36
NGC 3115	663	9.4	-24.0	135	122
NGC 3377	690	10.9	-22.8	92	84
NGC 4278	620	15.6	-23.8	172	145
NGC 4365	1243	23.3	-25.2	134	129
NGC 4473	2260	15.3	-23.8	63	60
NGC 4494	1342	16.6	-24.1	56	51
NGC 4649	1110	17.3	-25.5	153	145
NGC 5846	1712	24.2	-25.0	57	53

*Notes.* Column (1): galaxy name. Column (2): systemic velocity from Cappellari et al. (2011). Column (3): distance from Cappellari et al. (2011) based on surface brightness fluctuations (SBF) of Tonry et al. (2001). For NGC 4365, NGC 4473 and NGC 4649, Cappellari et al. (2011) used the SBF distance of Mei et al. (2007). Column (4):  $K$ -band absolute magnitude from Cappellari et al. (2011) calculated from the 2MASS (Skrutskie et al. 2006) total  $K$ -band apparent magnitude and the distance in Column (4), corrected for foreground extinction. Column (5): total number of GC spectra. Column (6): number of unique GCs with spectra. NGC 1407 and NGC 3115 are not part of the Cappellari et al. (2011) sample so their radial velocities were taken from NED while their distances are from Tonry et al. (2001) with the same  $-0.06$  mag distance moduli offset as Cappellari et al. (2011). We assume that NGC 1400 and NGC 1407 are part of the same group and use the mean of their distance moduli for NGC 1407. Their  $K$ -band absolute magnitude is also calculated from the 2MASS total  $K$  apparent magnitude and the distance in Column (4) in the same manner as Cappellari et al. (2011).

$\Delta\lambda \sim 1.5 \text{ \AA}$  and covers the Na I doublet to CaT wavelength region. In roughly half of the slits, H $\alpha$  is also covered. The DEIMOS data were reduced using the IDL SPEC2D pipeline (Cooper et al. 2012; Newman et al. 2013). We used the radial velocities measured by the sources of the spectra. These were measured using the IRAF<sup>2</sup> procedure FXCOR. Further details of the observations, data reduction and radial velocity measurements are given in Pota et al. (2013a).

For spectra observed in 2013 January, we found wavelength calibration issues at the bluest wavelengths ( $\lambda \sim 6500 \text{ \AA}$ ) with offsets of up to  $\sim 10 \text{ \AA}$  between the observed wavelengths of sky emission lines and their true wavelengths. We used the sky emission lines present in the background spectra to correct the wavelength issues of these spectra. This resulted in  $0.1 \text{ \AA}$  rms differences between the skylines and the corrected spectra.

For 6 of the 10 galaxies, we used the  $gi$  photometry from Usher et al. (2012) which is based on the Subaru Suprime-Cam, *Hubble Space Telescope* Wide-Field Planetary Camera 2 and *Hubble Space Telescope* Advanced Camera for Surveys (ACS) photometry in Pota et al. (2013a). For NGC 4278, we used the ACS and Suprime-Cam photometry presented in Usher et al. (2013) using equations A1 and A2 of Usher et al. (2012) to convert their  $gz$  photometry into  $gi$ . For objects in NGC 2768 with ACS photometry of Kartha et al. (2014), we used equations A6 and A7 of Usher et al. (2012) to

convert their  $BI$  photometry into  $gi$ . For the remaining GCs in NGC 2768, we used the Suprime-Cam  $gi$  photometry of Kartha et al. (2014). For both NGC 4473 and NGC 4649, we used the Suprime-Cam  $gi$  photometry from Pota et al. (in preparation). For seven of the NGC 4649 GCs without Suprime-Cam imaging, we used ACS  $gz$  photometry from Strader et al. (2012) using equations A1 and A2 of Usher et al. (2012) to convert their  $gz$  photometry into  $gi$ . For two NGC 4649 GCs without either Suprime-Cam or ACS photometry, we used CFHT MegaCam photometry from Pota et al. (in preparation) and used the photometric transformations provided on the MegaPipe website<sup>3</sup> to transform from the MegaCam filter system to the Sloan digital sky survey filter system used for the Suprime-Cam photometry. Thus, all the GCs in this study have original or transformed  $gi$  photometry.

### 3 ANALYSIS

For those spectra without CaT measurements from Usher et al. (2012) or Usher et al. (2013), we used the method of Usher et al. (2012) to measure the strengths of the CaT index. To reduce the effects of the strong sky line residuals found in the CaT spectral region, this method uses the technique of Foster et al. (2010) to mask the sky line regions before fitting a linear combination of stellar templates to the observed spectra using the PPIX pixel fitting code (Cappellari & Emsellem 2004). The fitted spectra are then normalized and the strength of the CaT lines measured on the normalized spectra. A Monte Carlo resampling technique is used to estimate 68 per cent confidence intervals for each CaT measurement. In Usher et al. (2012) when a single mask was observed on multiple nights, the spectra were combined from different nights into one spectrum before measuring the CaT. In this work, we considered the spectra from the same mask but different nights separately to allow telluric absorption to be corrected.

We excluded from our analysis those spectra with average S/N less than  $8 \text{ \AA}^{-1}$  in the wavelength range of 8400 to 8500  $\text{\AA}$ . Each spectrum was inspected and spectra with possible contamination from the host galaxy starlight or poor sky subtraction were excluded from our analysis. Out of the 1028 GCs in this study, we only have size information for only 391 (NGC 3115 from Jennings et al. 2014; NGC 4278 sizes from Usher et al. 2013; NGC 4365 from Blom, Spitler & Forbes 2012; NGC 4473 from Jordán et al. 2009; NGC 4649 from Strader et al. 2012). 10 of these have half-light radii larger than 10 pc, making them ultracompact dwarfs (UCDs) according to the definition of Brodie et al. (2011). Given that size information is not available for most of the GCs in this study, we did not distinguish between GCs and UCDs in our analysis. However, we did exclude the UCD M60-UCD1 (Strader et al. 2013) as it has a noticeably larger velocity dispersion ( $68 \text{ km s}^{-1}$ ) compared to the other GCs and UCDs in this study ( $<30 \text{ km s}^{-1}$ ).

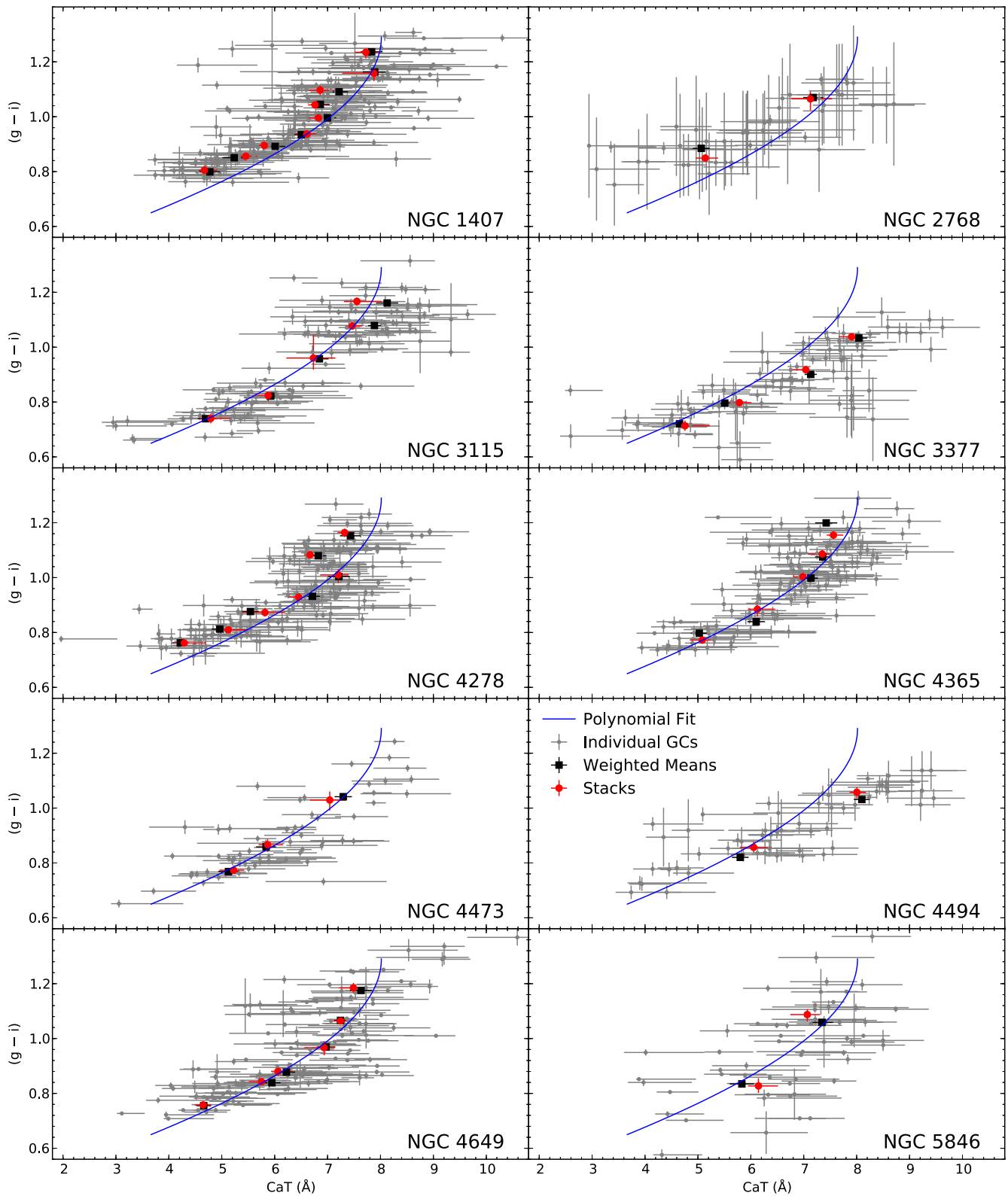
In Fig. 1, we plot the observed colours and measured CaT values of the individual GC spectra. To compare the relationship between colour and CaT between galaxies, we fit a cubic polynomial between colour and CaT to all GCs except those in NGC 2768 and NGC 5846, which both suffer from lower quality photometry:

$$\text{CaT} = -2.30(g - i)^3 - 3.32(g - i)^2 + 19.96(g - i) - 7.28. \quad (1)$$

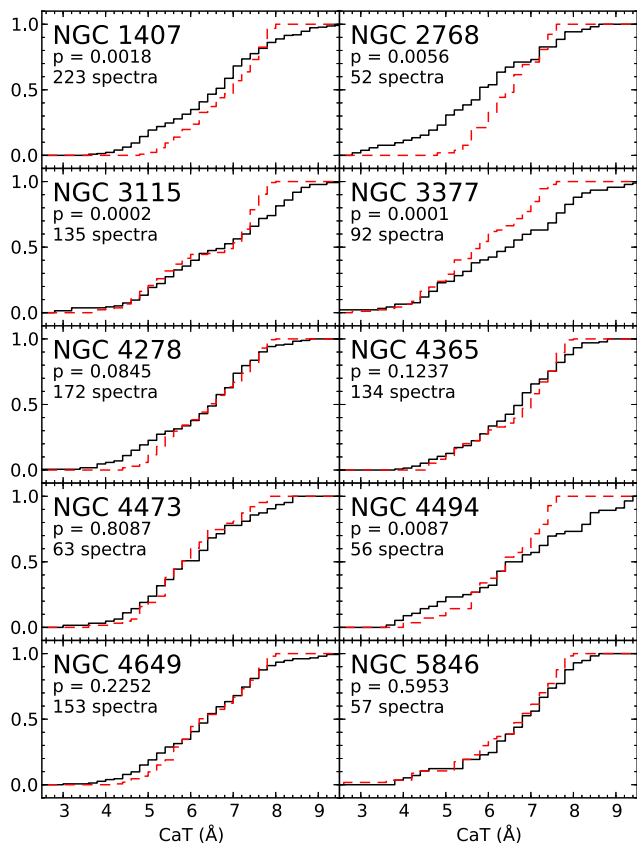
This fit shows good agreement with the observations in NGC 3115, NGC 4365, NGC 4473 and NGC 4649. In NGC 1407 and NGC 4278, the fit predicts CaT values higher than observed at blue colours

<sup>2</sup> IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

<sup>3</sup> <http://www3.cadc-ccda.hia-ihp.nrc-cnrc.gc.ca/en/megapipe/docs/filt.html>



**Figure 1.** CaT–colour diagrams for each of the ten galaxies in this study. The grey points represent the CaT measurements from individual spectra. The black squares represent the weighted means of the individual galaxies spectra using the same binning as used for stacking. The red circles show the measurements from the stacked spectra. The blue lines show a cubic polynomial fit to the individual spectra in all galaxies combined except NGC 2768 and NGC 5846. Note that the blue line is the same in all panels. In general, the stacked spectra and the mean values agree. Note that the shape of the CaT–colour relation is near linear in NGC 3377 and NGC 4494 while in NGC 1407 and NGC 4278 the blue GCs have lower CaT strengths than predicted by the fitted polynomial. The CaT–colour relation varies from galaxy to galaxy.



**Figure 2.** Cumulative CaT histograms for each galaxy. The solid black lines show the measured CaT distributions while the dashed red lines show the CaT distributions predicted by the colour–CaT relation given in equation (1). The Kolmogorov–Smirnov test probability that these are both drawn from the same distribution is printed for each galaxy along with the number of spectra. Although the CaT distributions of galaxies like NGC 4365 and NGC 4649 are well predicted by their colour distributions, in NGC 1407 the colours overpredict the CaT values for blue colours, while in NGC 3377 and NGC 4494 the colours underpredict the CaT values for red colours. This is further evidence that the CaT–colour relation varies from galaxy to galaxy.

while in NGC 3377 and NGC 4494 the fit predicts CaT values lower than observed at red colours. These results are the same as those found by Usher et al. (2012).

In Fig. 2, we compare cumulative CaT histograms of the observed CaT distribution with the CaT distribution predicted from the colour distribution using equation (1). We ran Kolmogorov–Smirnov tests to quantify the difference between the two CaT distributions. Some galaxies such as NGC 4365 and NGC 4649 show agreement between the measured CaT values and those predicted from colours, while other galaxies show disagreement with the colours, predicting too few CaT weak GCs in NGC 1407 and too few CaT strong GCs in NGC 3377. As in Usher et al. (2012), we see evidence that the colour–CaT relation varies from galaxy to galaxy.

### 3.1 Stacking the spectra

We studied two spectral regions. The first, which covers from 8100 to 8900 Å, contains the CaT at 8498, 8542 and 8662 Å, the sodium doublet at 8183 and 8195 Å (Na82), the magnesium line at 8807 Å and a whole host of weaker metal lines. The second, which covers from 6502–6745 Å, contains the H $\alpha$  line at 6563 Å. In the blue region, we only used spectra with coverage bluewards of 6540 Å

in the rest frame so that only spectra with observed H $\alpha$  lines were included. Since the grating setup we used only covers H $\alpha$  in only approximately 60 per cent of the slits (674 of 1137), the S/N of the blue regions of the stacked spectra are lower than those of the red region of the stacks.

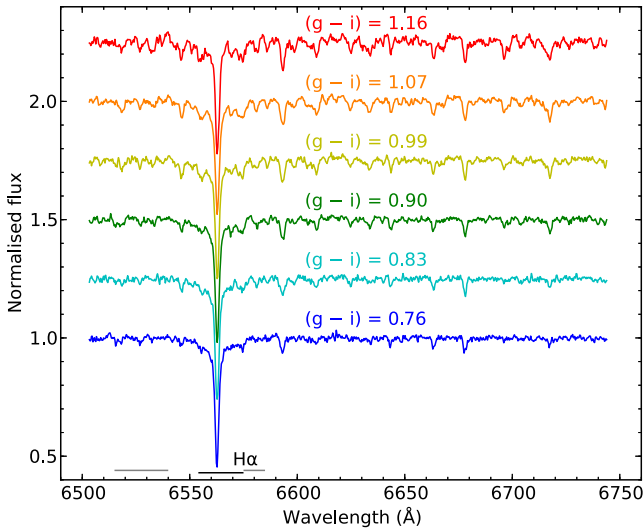
The spectral region from 8128 to 8348 Å, which contains the sodium doublet, is affected by telluric absorption lines. To remove the effect of these lines, we observed the spectroscopic standard star BD +28 4211 on 2012 October 20 using DEIMOS in long-slit mode with the same spectroscopic setup as the GC spectra for a total of 250 s. Likewise, the long-slit data were also reduced using IDL SPEC2D pipeline. For each observed mask, we stacked all GCs from that mask using the observed wavelengths rather than shifting them to the rest wavelengths. We then scaled the normalized telluric absorption spectrum to match the telluric lines in the stacked spectra. Lastly, we divided each observed spectrum by the scaled telluric spectrum to remove the effects of the telluric lines. Unfortunately, NGC 1407’s radial velocity places the Na82 feature at the same observed wavelength as the strongest telluric absorption in this wavelength range. For four masks in NGC 1407 with the strongest telluric absorption, we were unable to satisfactorily remove the absorption. The variances of these spectra in the affected wavelength range were multiplied by a hundred in order to down-weight them in the stacking procedure. NGC 5846’s radial velocity also places the Na82 feature in the same wavelength range as strong telluric lines but we did not down-weight its spectra as the strength of the telluric absorption was not as strong as in the NGC 1407 observations.

Before stacking, the spectra were shifted to the rest frame and re-dispersed to a common linear wavelength scale of 0.33 Å per pixel. We do not smooth the spectra to a common dispersion before stacking. We continuum normalized each red spectrum by fitting and dividing by a linear combination of Chebyshev polynomials of the order of 0–8 while each blue spectrum was normalized using Chebyshev polynomials of the order of 0–5. When fitting the polynomials we masked wavelength ranges identified from the Vazdekis et al. (2012) stellar population models as containing strong absorption lines. Next, we grouped the spectra and stacked the spectra in each group. At each pixel, the flux from each spectrum in a group was weighted by its variance and summed. For each group, we calculated a mean colour and a mean magnitude by using the pixel weights in the 8400 to 8500 Å range.

First, we stacked the spectra from all galaxies in six bins of equal size by colour as seen in Figs 3 and 4. For the red spectral region, these stacks have S/N of 248–320 Å<sup>−1</sup> while in the blue they have S/N of 159–239 Å<sup>−1</sup>. Secondly, for each galaxy, the spectra were grouped by  $(g - i)$  colour in bins of  $\sim 25$  spectra. For these stacks, the S/N in the red region is between 68 and 155 Å<sup>−1</sup> with a median of 94 Å<sup>−1</sup> while in the blue it is between 32 and 140 Å<sup>−1</sup> with a median of 70 Å<sup>−1</sup>.

### 3.2 Spectral feature measurements

To compare the stacked spectra in a quantitative manner we measured several spectral indices from the stacked spectra. As noted by Conroy & van Dokkum (2012a), spectral indices suffer from several shortcomings. The strength of an absorption feature is measured relative to a pseudo-continuum. The level of the pseudo-continuum is also affected by absorption lines so the variations in the measured index can be due to the feature of interest, the pseudo-continuum or both. Spectral indices rarely contain absorption lines from just one species, complicating their interpretation. Since a spectral index collapses all information about a feature into a single value, useful



**Figure 3.** Normalized and offset stacks of all 674 GC spectra in the  $H\alpha$  region. The spectra have been grouped and colour coded by their  $(g - i)$  colour. The horizontal lines show the definition of the  $H\alpha$  index. The black line at the bottom shows the feature passband while the grey lines show the pseudo-continuum passbands. Most noticeable is the strengthening of the weak metal lines (e.g. at 6593 and 6717 Å) with redder colours. The  $H\alpha$  line itself becomes slightly weaker with redder colours.

information from the shape of a spectral feature is lost. However, spectral indices do allow for spectra to be compared in a model-independent manner.

We measured the strength of the CaT feature in the stacked spectra using the technique in Usher et al. (2012); see above for more details. As part of the CaT measurement process, the velocity dispersion of each spectrum is fitted. We measured the strength of the Mg 1 line at 8807 Å using the index definition of Cenarro et al. (2009), here denoted Mg88. We defined a weak metal line index, Fe86, consisting of Fe and Ti lines in the spectral region of the CaT. The index was defined by identifying Fe and Ti lines that changed with metallicity in the Vazdekis et al. (2012) models in the CaT spectral

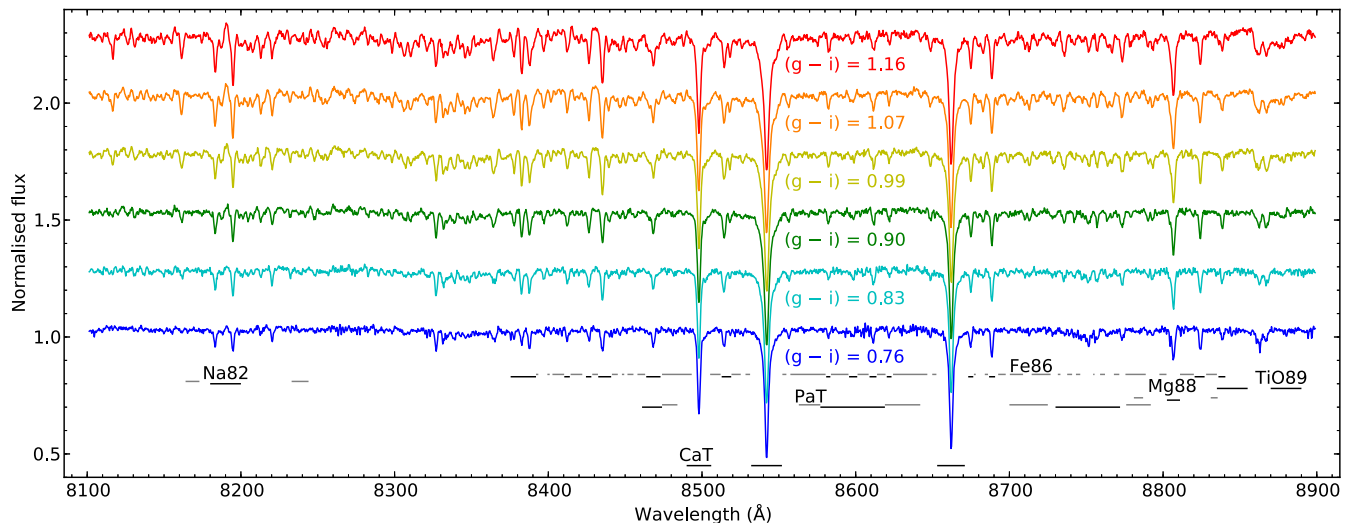
region and by examining the stacks of all GCs plotted in Fig. 4. A subset of these lines was studied by Foster et al. (2010). We used the  $H\alpha_A$  index of Nelan et al. (2005) to measure the strength of  $H\alpha$  (here denoted  $H\alpha$ ). To measure the strength of the Na I doublet at 8183 and 8195 Å, we used the Na I 8200 index of Vazdekis et al. (2012), here denoted Na82. We used the PaT index of Cenarro et al. (2001) to measure the strength of the hydrogen Paschen lines in the region of the CaT. To measure the strength of the TiO band head at 8859 Å we used the flux ratio defined by Conroy & van Dokkum (2012a), here denoted TiO89. Except for the CaT strengths and the TiO89 flux ratios, we used the method of Cenarro et al. (2001) to calculate the index values. The index definitions are given in Table 2 and plotted in Figs 3 and 4. These measurements are given in Table A1 and will be discussed in Sections 4.1, 4.2 and 4.3.

To estimate the uncertainty of the measured spectral features, as well as the colours and magnitudes of the stacks, we used bootstrapping. For each galaxy, 256 random samples of the GC spectra from that galaxy were created with replacement with sizes equal to the number of GC spectra in that galaxy. The stacking process and line index measurements were repeated for each of these samples. For the colour, the magnitude and each of the measured spectral features, the distribution of that property in the bootstrap samples was smoothed with a Gaussian kernel. For each of these distributions, a 68.3 per cent confidence interval was calculated by integrating the likelihood function over all values where the likelihood function is larger than some constant value and varying the constant value until the integral of the likelihood equals 0.683. The limits of the integration are then the confidence interval (Andrae 2010).

In Fig. 1, we compare the colours and CaT values of the stacks with those of individual GC spectra. We also calculated the variance weighted means of colour and CaT in the same bins as the stacks. We see good agreement between the stack colours and CaT values and the means of the individual measurements in the same bin.

### 3.3 Trends with magnitude

As can be seen in the colour–magnitude diagram in Fig. 5, the absolute magnitudes of the GCs in this study range from  $M_i = -13.1$

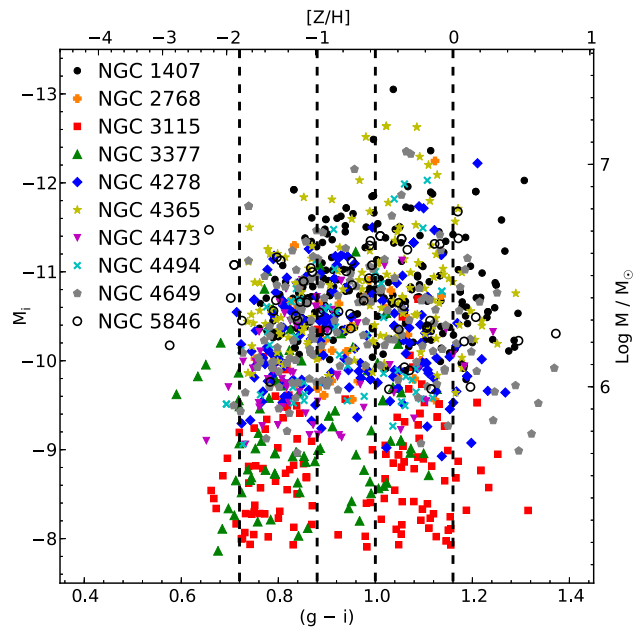


**Figure 4.** Normalized and offset stacks of all 1137 GC spectra in the region around Na82 and CaT features. The spectra have been grouped and colour coded by their  $(g - i)$  colour. The horizontal lines show the index definitions. The black line at the bottom shows the feature passband while the grey lines show the pseudo-continuum passbands. Going from the bluest spectra to the reddest the strengths of the metal lines increase. At the reddest colours, TiO molecular features (such as at  $\sim 8190$ ,  $\sim 8430$  and  $\sim 8860$  Å) appear.

**Table 2.** Spectral index definitions.

Name	Feature passbands (Å)	Pseudo-continuum passbands (Å)	Reference
(1)	(2)	(3)	(4)
CaT	8490.0–8506.0	Polynomial normalization	U12
	8532.0–8552.0		
	8653.0–8671.0		
Mg88	8802.5–8811.0	8781.0–8787.9	C09
		8831.0–8835.5	
Fe86	8375.5–8382.0	8392.0–8393.5	This work
	8410.4–8414.0	8399.4–8400.9	
	8424.5–8428.0	8402.7–8410.3	
	8432.5–8440.0	8414.5–8422.1	
	8463.7–8473.0	8428.6–8432.3	
	8512.8–8519.0	8441.4–8445.2	
	8580.8–8583.5	8447.9–8449.4	
	8595.7–8601.0	8451.5–8455.4	
	8609.0–8613.5	8458.0–8463.0	
	8620.2–8623.3	8474.0–8493.3	
	8673.2–8676.5	8505.3–8512.1	
	8686.8–8690.7	8519.2–8525.2	
	8820.5–8827.0	8528.3–8531.3	
	8836.0–8840.6	8552.3–8554.9	
		8557.5–8580.4	
		8583.9–8595.3	
		8601.2–8608.4	
		8613.9–8619.4	
		8624.3–8646.6	
		8649.8–8652.5	
	8676.9–8678.1		
	8684.0–8686.1		
	8692.7–8697.6		
	8700.3–8708.9		
	8714.5–8726.8		
	8731.5–8733.2		
	8737.6–8740.8		
	8743.3–8746.1		
	8754.5–8755.4		
	8759.0–8762.2		
	8768.0–8771.5		
	8775.5–8788.7		
	8797.6–8802.2		
	8811.0–8820.0		
	8828.0–8835.0		
	8835.0–8855.0		
TiO89		8870.0–8890.0	CvD12
Na82	8180.0–8200.0	8164.0–8173.0	V12
		8233.0–8244.0	
H $\alpha$	6554.0–6575.0	6515.0–6540.0	N05
		6575.0–6585.0	
PaT	8461.0–8474.0	8474.0–8484.0	C01
	8577.0–8619.0	8563.0–8577.0	
	8730.0–8772.0	8619.0–8642.0	
		8700.0–8725.0	

Notes. Column (1): index name. Column (2): feature passbands in Å. Column (3): pseudo-continuum passbands in Å. Column (4): reference for index definition. U12: Usher et al. (2012), C09: Cenarro et al. (2009), CvD12: Conroy & van Dokkum (2012a), V12: Vazdekis et al. (2012), N05: Nelan et al. (2005), C01: Cenarro et al. (2001).



**Figure 5.** Colour–magnitude diagram for all the GC spectra used in this study.  $i$ -band absolute magnitude is plotted versus  $(g - i)$  colour. The dashed vertical lines show the colour cuts used when stacking GCs by magnitude. On the top axis, the corresponding metallicities based on the colour–metallicity relation of Usher et al. (2012, their equation 10) are shown. For illustrative purposes, the mass corresponding to the absolute magnitude and a constant mass-to-light ratio of 2 is shown on the right. More distant and luminous galaxies such as NGC 1407 and NGC 4365 contribute GCs predominantly brighter than  $M_i \sim -10$ , while closer galaxies such as NGC 3115 and NGC 3377 contribute fainter GCs. Note that the GC selection varies galaxy to galaxy so that the colour distribution of GCs with CaT measurements may not be representative of the overall GC colour distribution within that galaxy. A colour–magnitude relation can be seen for the blue GCs ( $g - i < 0.94$ , a ‘blue tilt’) which is much stronger for GCs brighter than  $M_i = -11$ . Almost all GCs brighter than  $M_i = -12$  are red.

to  $M_i = -7.9$ . Assuming a mass-to-light ratio of 2 (see below), this corresponds to a mass range of  $2.2 \times 10^7$ – $1.8 \times 10^5 M_\odot$ . The bright GCs tend to come from different galaxies than the faint GCs. This makes it difficult to tell if any observed trends with magnitude are really due to magnitude or are due to differences between galaxies (i.e. the variations in colour–CaT relations seen between galaxies, Fig. 1). To attempt to disentangle galaxy to galaxy variation from trends with luminosity, we split the galaxies into two groups by the shape of their colour–CaT relation and stacked the spectra by magnitude within each group. We included NGC 3115, NGC 3377, NGC 4473 and NGC 4494 in one group and NGC 1407, NGC 4278, NGC 4365 and NGC 4649 in another. The galaxies in the first group show bluer colours at low and at high CaT strengths compared with the second group (see Fig. 1). Since both the GC and galaxy luminosities are higher on average in the second group than in the first, we will refer to the groups as the faint and bright groups, respectively.

We combined GC spectra from each group in two colour ranges: blue with colours in the range  $0.72 < (g - i) < 0.88$  and red with colours in the range  $1.00 < (g - i) < 1.16$ . The limits of the colour ranges can be seen in the Fig. 5. We stacked the faint group spectra in four bins by magnitude and the bright group in six bins by magnitude. The S/N of faint group stacks ranged from 67 to  $198 \text{ \AA}^{-1}$  with a median of  $97 \text{ \AA}^{-1}$  in red and from 44 to  $168 \text{ \AA}^{-1}$  with a median of  $73 \text{ \AA}^{-1}$  in the blue while the bright group

stacks have a red S/N between 69 and 244  $\text{\AA}^{-1}$  with a median of 99  $\text{\AA}^{-1}$  and a blue S/N between 39 and 161  $\text{\AA}^{-1}$  with a median of 76  $\text{\AA}^{-1}$ . We measured the colours, magnitudes and line indices strengths as before. 1024 bootstrap iterations were used to estimate the uncertainties on these measurements. The measured colours, magnitudes and spectral indices in Table A2. Using equation (5), the mean CaT value of the blue stacks corresponds to  $[Z/H] = -1.21$  while the CaT value of the red stacks corresponds to  $[Z/H] = -0.45$ .

We used the Conroy et al. (2009) simple stellar population models with an age of 12.6 Gyr and a Kroupa (2001) IMF to estimate the *i*-band mass-to-light ratio. The model mass-to-light ratio increases from 1.8 at  $[\text{Fe}/\text{H}] = -2.0$  to 2.3 at  $[\text{Fe}/\text{H}] = -0.5$  before rising sharply to 3.7 at  $[\text{Fe}/\text{H}] = 0.2$ . Beyond age and metallicity, the IMF and the horizontal branch morphology also affects the mass-to-light ratios (Conroy et al. 2009). In addition to stellar population effects, dynamical evolution preferentially removes low-mass stars from GCs (e.g. Hénon 1969; Baumgardt & Makino 2003). The relative fraction of low-mass stars lost is dependent on the mass of the GC and on the IMF (Goudfrooij & Kruijssen 2014). Strader, Caldwell & Seth (2011a) found that although the mass-to-light ratios of metal-poor GCs in M31 are in agreement with the predictions of simple stellar population models with a Kroupa (2001) IMF from Conroy et al. (2009), the metal-rich GCs in M31 have lower than predicted mass-to-light ratios. Due to the uncertainty regarding the correct IMF to use, for illustrative purposes we adopt a mass-to-light ratio of 2 for all metallicities and study trends with luminosity not mass. The stacks range in luminosity from  $4.7 \times 10^6$  to  $1.5 \times 10^5 L_{\odot}$ .

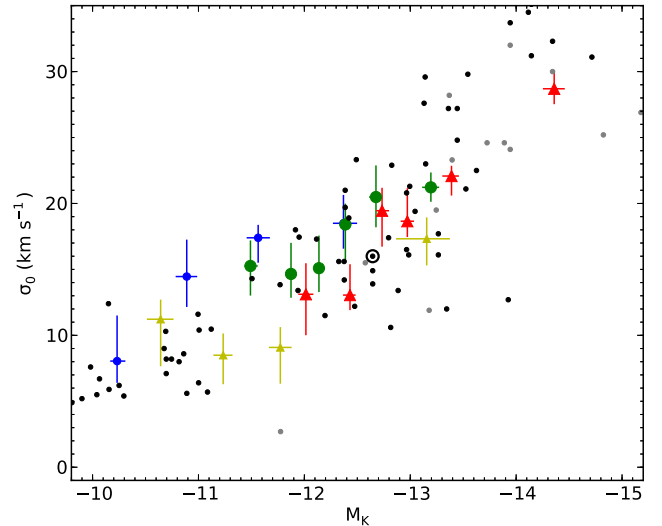
As a check on the stacking procedure, in Fig. 6 we compare the velocity dispersions measured from the stacked spectra by PPF to the Forbes et al. (2008) compilation of GC and UCD kinematics. To convert our *i*-band absolute magnitudes into *K*-band magnitudes, we use the Conroy et al. (2009) simple stellar population models with an age of 12.6 Gyr and a Kroupa (2001) IMF to find the following relation between  $(i - K)$  and  $(g - i)$ :

$$(i - K) = 1.364 \times (g - i) + 0.756. \quad (2)$$

This relation has an rms of 0.034 mag. Since the GCs in this study are unresolved, we follow Forbes et al. (2008) and multiply the measured velocity dispersions by 10 per cent to estimate the central velocity dispersions. As can be seen in Fig. 6, the velocity dispersion trends of the stacks agree with the trend for GCs and UCDs from Forbes et al. (2008), suggesting that our stacking procedure is reliable. The metal-poor stacked spectra appear to have higher velocity dispersions at fixed absolute magnitudes compared to the metal-rich stacks. This is in agreement with the results of Strader et al. (2011a) who found that metal-poor GCs have significantly higher *K*-band mass-to-light ratios compared to metal-rich GCs. We plan to further investigate GC velocity dispersions and mass-to-light ratios in a future paper.

To investigate trends with magnitude, we fit linear relations between the absolute magnitude and the various spectral features as well as the colours. We used 1024 bootstrap iterations to estimate 68 per cent confidence intervals for the slope and intercept as well as the probability that the slope is non-zero. We list the fitted slopes and intercepts as well as the probability that the fitted slope is less than zero in Table B1. We use the bootstrap probability that the fitted slope is less than zero being greater than 0.95 for negative slopes or less than 0.05 for positive slopes as criteria for a significant relation. Significant relations are bold in Table B1.

To allow the two galaxy groups to be compared at fixed luminosity, we used the fitted relations with magnitude to calculate the



**Figure 6.** Central velocity dispersion versus absolute *K*-band magnitude. The stacks of blue GC spectra in the faint galaxy group are plotted as small blue circles, blue GC spectra in the bright galaxy group as large green circles, red GC spectra in the faint group as yellow small triangles and red GC spectra in the bright group as red large triangles. Black points represent GCs from Forbes et al. (2008) (half-light radii  $< 10$  pc) while grey points UCDs (half-light radii  $> 10$  pc) from the same source. The circled point represent  $\omega$  Centauri. Equation (2) was used to transform the stacks' mean *i*-band magnitudes into *K*-band magnitudes. In line with Forbes et al. (2008), the velocity dispersions measured from the stacked spectra was increased by 10 per cent to estimate the central velocity dispersion. The stacks' velocity dispersion trends agree with that from Forbes et al. (2008), suggesting that our stacking procedure is reliable. The data suggest that metal-poor (blue) stacks have higher velocity dispersions at fixed magnitude in line with the results of Strader et al. (2011a) who found that metal-poor GCs have significantly higher *K*-band mass-to-light ratios compared to metal-rich GCs.

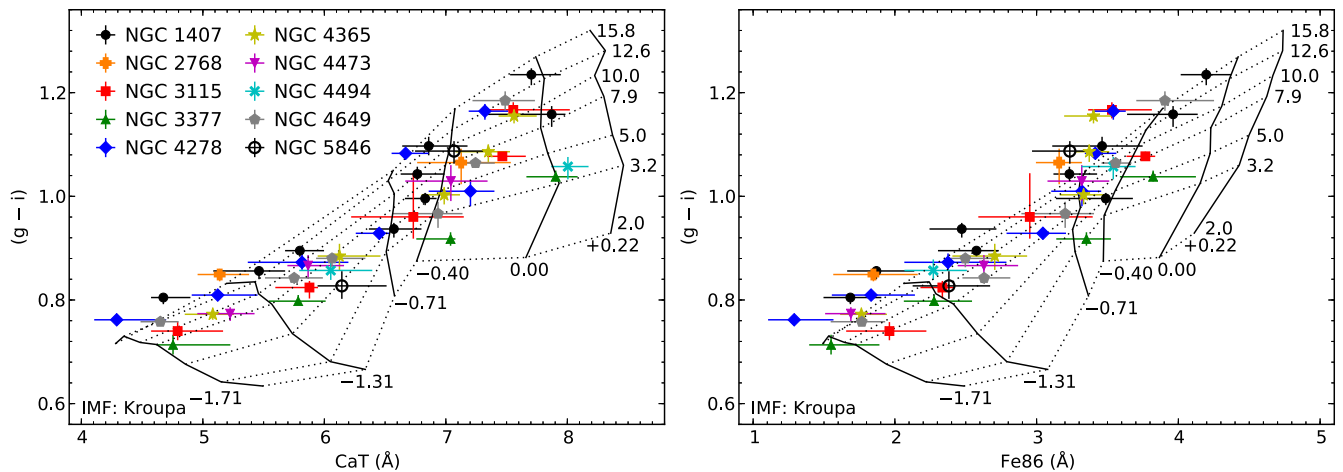
value of each spectral feature and colour at  $M_i = 10$ . We also used the bootstrap samples to calculate the probability that difference between each feature at this absolute magnitude is greater than zero. As before we consider probabilities above 0.95 for positive differences and below 0.05 for negative differences to be significant. The values of each feature at  $M_i = 10$  and the probabilities that the differences between the two groups are greater than zero is given in B1. We will discuss the colours and index strengths of the GC spectra stacked by magnitude in Sections 4.3 and 4.4.

## 4 DISCUSSION

### 4.1 Comparison with stellar population models

To gain some understanding on how these indices and the  $(g - i)$  colour are predicted to respond to changes in metallicity, age and the IMF, we analysed the simple stellar population models of Vazdekis et al. (2012). The spectral resolution of these models (1.5  $\text{\AA}$  in the CaT spectral region, 2.5  $\text{\AA}$  in the regions of  $H\alpha$  and Na82) closely matches that of the DEIMOS spectra and the models span a wide range of metallicities, unlike the more advanced models of Conroy & van Dokkum (2012a), as well as a range of IMFs. Additionally, the Vazdekis et al. (2012) model spectral energy distributions are publicly available. Usher et al. (2012) found that the Vazdekis et al. (2012) models predicted a CaT- $[Z/H]$  relation in good agreement with observations.





**Figure 7.** Colour versus metal line indices for GC spectra stacked by colour. On the left,  $(g - i)$  colour is plotted as a function of the CaT index, while on the right,  $(g - i)$  colour is plotted as a function of the Fe86 index. In both panels, the Vazdekis et al. (2012) model predictions for a Kroupa (2001) IMF and varying age and metallicity are overplotted. For each model grid, solid lines are lines of constant metallicity; dotted lines are constant age. The points are colour coded by their host galaxy using the same colours and shapes as in Fig. 5. Assuming universally old ages, the colour–CaT plot shows poor agreement between the observations and the models with the models predicting redder colours than observed at intermediate to strong CaT values. However, the colour–Fe86 plot shows good agreement between the models and the stacked spectra. At high CaT values, NGC 4494 is an outlier in the colour–CaT plot but is consistent with the other galaxies in the colour–Fe86 plot.

The Vazdekis et al. (2012) models use the scaled solar isochrones from Girardi et al. (2000) with use the stellar libraries of Cenarro et al. (2001) in the CaT region (8350 to 9020 Å), MILES (Sánchez-Blázquez et al. 2006) from 3525 to 7500 Å and Indo-US (Valdes et al. 2004) from 7500 to 8350 Å. These spectral libraries are dominated by field stars but contain a number of open cluster and metal-poor GC stars. Since the abundance pattern of stars in the solar neighbourhood is not that of GCs in the Milky Way, the models likely do not accurately predict the strengths of some spectral features. For the CaT, Mg88, Fe86, PaT and TiO89 indices, we used the Vazdekis et al. (2012) models based on the stellar library of Cenarro et al. (2001). For the  $H\alpha$  and Na82 indices as well as the  $(g - i)$  colour, we used the Vazdekis et al. (2012) models based on the MIUSCAT stellar library (Vazdekis et al. 2012). To minimize systematic effects due to flat fielding, we applied the same normalization that we applied to the observed spectra to the model spectral energy distributions before measuring the spectral features in an identical fashion. Due to the similarity of the model spectral resolution to that of the observed spectra, we did not consider what effect spectral resolution has on the index measurement.

We compare the colours and spectral indices of the stacked spectra with the Vazdekis et al. (2012) models in Figs 7–9. Although the models predict the general colour and index–CaT trends observed, the models disagree with the observations in several details. In Fig. 7, we plot colour as functions of CaT and Fe86 index strengths. Assuming that all GCs have similar old ages, the Vazdekis et al. (2012) models predict redder colours for GCs at intermediate CaT strengths than are observed. Some of the disagreement between the models and the observed values may be due to differences in  $\alpha$ -element abundances between the models and the observed GCs. The observed colour–Fe86 relation generally agrees with the model predictions. As seen the top-right panel of Fig. 8, the observed colour–CaT relation suggests a bottom light IMF.

As seen in the middle row of Fig. 8, the predicted relationship between the Na82 and CaT indices is harder to assess due to the likely non-solar [Na/Fe] abundance in the GCs and the possible systematics due to correcting the telluric absorption present in this

wavelength region. Additionally, Vazdekis et al. (2012) state that their models for the Na82 wavelength region are unreliable for metallicities below  $[Z/H] = -0.71$ . At high CaT strengths, the majority of the stacked spectra suggest a Milky Way like IMF (i.e. Kroupa 2001, like). If the stacked spectra show enhanced [Na/Fe] like Milky Way GCs (e.g. Roediger et al. 2014), the Na82 index should be stronger (Conroy & van Dokkum 2012a), strengthening the evidence for bottom light IMF.

The Fe86 index strengths show a tight relation with the CaT index strengths that is parallel but offset from the predictions of Vazdekis et al. (2012). Interestingly, the models of Conroy & van Dokkum (2012a) underpredict the observed strength of the CaT in galaxy spectra (Conroy & van Dokkum 2012b) by a similar amount ( $\sim 0.5$  Å). Fitting the Fe86 index strength as a function of CaT index strength, we find

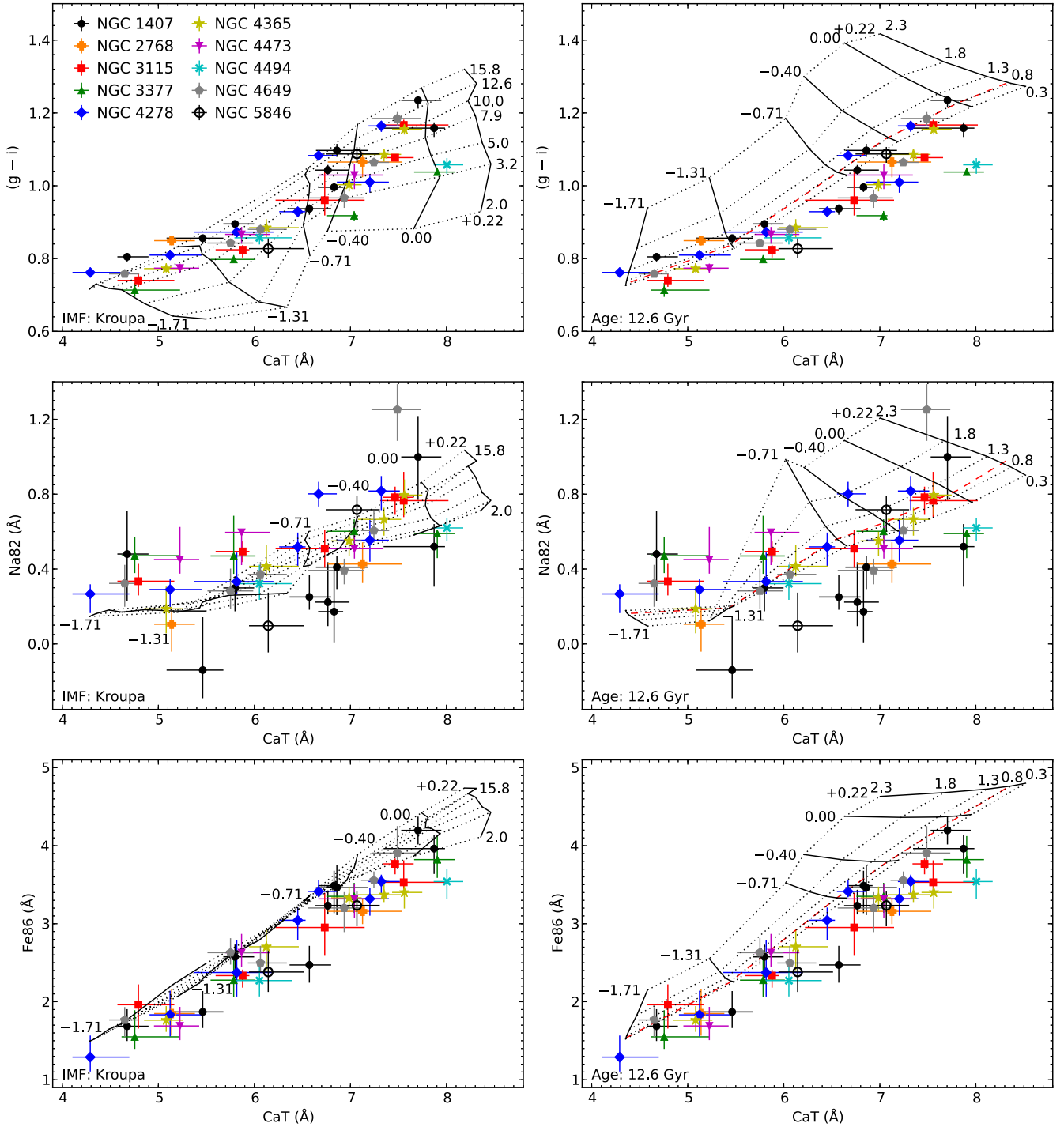
$$\text{Fe86} = (0.724 \pm 0.032) \times \text{CaT} + (-1.802 \pm 0.209). \quad (3)$$

This fit has an rms of 0.212 Å and a  $\chi^2$  value of 34.7 for 43 degrees of freedom. Foster et al. (2010) observed a similar relation between a subset of the lines in the Fe86 index and the CaT. However, they were only able to measure the strength of the weak metal lines on the spectra fitted as part of their CaT measurement process. The relation between the Mg88 and CaT indices is also linear but shows more scatter than the Fe86–CaT index relation. Fitting the Mg88 index strengths, we find

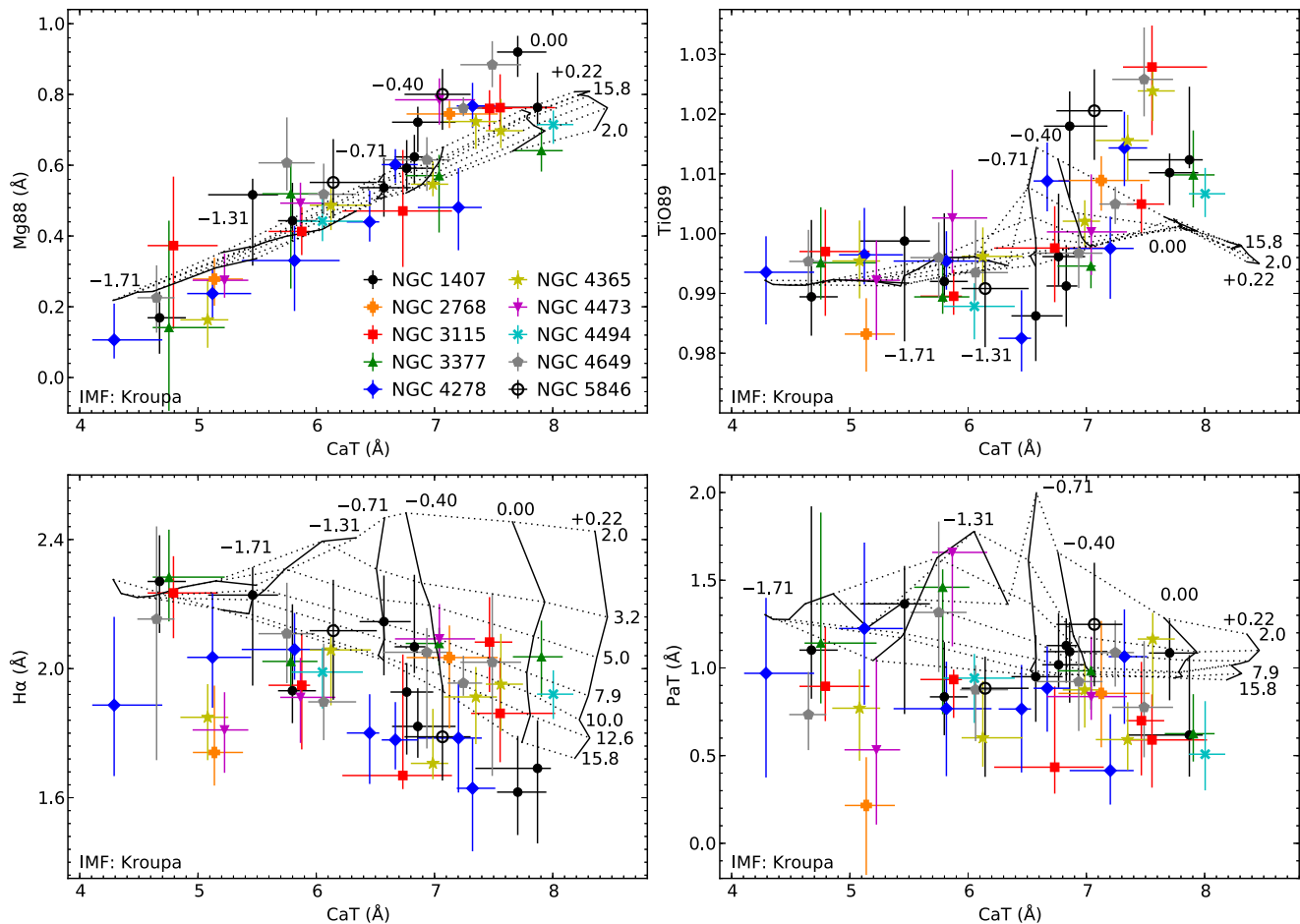
$$\text{Mg88} = (0.184 \pm 0.014) \times \text{CaT} + (-0.643 \pm 0.091). \quad (4)$$

This fit has an rms of 0.092 Å and a  $\chi^2$  value of 45.3 for 43 degrees of freedom. Although most Mg88 strengths agree with the Vazdekis et al. (2012) models, at low CaT values the Mg88 strengths are lower than the models while at high CaT values, the Mg88 values are higher than the models. These differences between the observations and the Vazdekis et al. (2012) models are likely due to differences in the abundance patterns of the stellar libraries used in the models compared to the GC abundance patterns.

The strength of the TiO89 molecular feature appears flat at low CaT values but increases with CaT strength at higher CaT values.



**Figure 8.** Colours and line indices for GC spectra stacked by colour. In the left-hand column, the Vazdekis et al. (2012) model predictions for a Kroupa (2001) IMF and varying age and metallicity are overplotted. In the right-hand column, the Vazdekis et al. (2012) model predictions for varying IMF slope and metallicity at a fixed age of 12.6 Gyr are overplotted in black. A slope of 1.3 corresponds to a Salpeter (1955) IMF. The 12.6 Gyr Kroupa (2001) IMF model is plotted with a red dashed line. For each model grid, solid lines are lines of constant metallicity; dotted lines are constant age or IMF slope. Top:  $(g-i)$  colour as a function of the CaT index. At low CaT values NGC 3115, NGC 3377 and NGC 4473 show bluer colours than the other galaxies while NGC 3377 and NGC 4494 show bluer colours at high CaT values. Assuming universally old ages, the Vazdekis et al. (2012) models do not predict the observed relation between CaT and colour. The colours and CaT values do suggest a bottom light IMF. Middle: the Na82 sodium index as a function of the CaT index. The Na82 values are likely affected by systematics due in the removal of telluric absorption lines. At low CaT values NGC 3115, NGC 3377 and NGC 4473 show higher Na82 strengths than other galaxies and while at high CaT values both NGC 3377 and NGC 4494 appear to have lower Na82 values. At high metallicities, the Na82 and the CaT respond in opposite directions to the slope of the IMF. However, the Na82 is affected by the Na abundance which is known to be non-solar in GCs. At lower metallicities, the Vazdekis et al. (2012) models are less reliable. Bottom: the Fe86 weak metal line index as a function of the CaT index. The similar relations between the indices for different galaxies suggest that both indices are measuring metallicity. Unlike the CaT, Fe86 is insensitive to the IMF at high metallicities.



**Figure 9.** Line indices for GC spectra stacked by colour. The points are colour coded by their host galaxy using the same colours and shapes as in Fig. 8. In each plot, the Vazdekis et al. (2012) model predictions for a Kroupa (2001) IMF and varying age and metallicity are overplotted. For each model grid, solid lines are lines of constant metallicity; dotted lines are of constant age. Top left: the Mg88 magnesium index as a function of CaT. All galaxies show similar Mg88–CaT trends, suggesting that the  $[Mg/Fe]$ – $[Z/H]$  relation is similar for all galaxies. However, the observed relation is tilted compared to the Vazdekis et al. (2012) models. Top right: the TiO89 TiO index versus the CaT index. At low CaT values, the measured strength of this molecular feature is independent of the CaT value. At high CaT values the strength increases with the CaT values. The Vazdekis et al. (2012) models poorly predict the strength of TiO89 at high metallicities. Bottom left: the  $H\alpha$  index versus the CaT index. Overall  $H\alpha$  strength decreases with CaT strength in line with the model predictions but there is a wide scatter in  $H\alpha$  strengths. Many of the stacks show weaker  $H\alpha$  strengths than the predictions of the oldest models. NGC 3377 and NGC 4494, which show the bluest colours at high CaT, show the strongest  $H\alpha$  at high CaT, suggesting younger ages. Bottom right: the PaT Paschen line index versus the CaT index. Although the strengths of the Paschen lines show large scatter, the mean values are weaker than the oldest models.

Although the Vazdekis et al. (2012) models predict the behaviour of the TiO89 index at weak CaT strengths, the models underpredict this TiO index at stronger CaT values. This is likely due to the solar abundance pattern of the stellar library used by Vazdekis et al. (2012) at high metallicities as the TiO molecular bands are strongly sensitive to both the O and Ti abundances (Conroy et al. 2014). Since the TiO89 index is defined as a flux ratio, it is sensitive to differences in flux calibration. However, differences in the TiO89 index due to flux calibration should be minimized since we normalized the continuum in an identical fashion for both the observed spectra and the model spectra. The increase in strength of the TiO89 feature with metallicity is apparent in Fig. 4. Although there is a larger amount of scatter, the observed  $H\alpha$  values decrease with increasing CaT strength in line with the model predictions. Likewise the PaT values show a large amount of scatter but no trend with CaT, again in agreement with the models. The lack of significant PaT absorption in the stacked spectra allows us to rule out truly young ages for any of the stacked spectra. For both  $H\alpha$  and PaT, the mean index measurements are systematically weaker than

the oldest Vazdekis et al. (2012) model predictions. Foster et al. (2010) observed significant Paschen line strengths in their fitted spectra but not in the raw spectra. Usher et al. (2012) refit these spectra and observed no Paschen absorption in the fitted spectra. Since we do not observe Paschen line absorption in our stacks, the absorption observed by Foster et al. (2010) was likely an artefact of their fitting procedure. Combining the  $(g - i)$  colours as well as the CaT, Na82, Fe86 and  $H\alpha$  strengths with the Vazdekis et al. (2012) models, the stacked spectra suggest old ages and a bottom light IMF.

Vazdekis et al. (2012) state that their models are unreliable at young ages for model metallicities below  $[Z/H] = -0.40$  and that the models in the Na82 spectral region are unreliable for metallicities below  $[Z/H] = -0.71$  for any age. The inability to vary chemical abundances from the solar neighbourhood pattern or vary the horizontal branch morphology are limitations to using the Vazdekis et al. (2012) models to further interpret our data. Medium-resolution stellar population models with the ability to fully vary chemical abundances and to adopt different horizontal branch morphologies

would be necessary to recover more information from the DEIMOS spectra.

#### 4.2 Fe86-based metallicities

Although Usher et al. (2012) showed that the CaT can be used to derive metallicities as well as the traditionally employed Lick indices (Worthey et al. 1994), it is desirable to further check the reliability of the CaT as a metallicity indicator. As noted before, Kirby et al. (2008) used the weak metal lines in DEIMOS spectra of RGB stars to derive iron and  $\alpha$ -element abundances. Our Fe86 index measures some of the stronger lines fit by Kirby et al. (2008) and as seen in the lower-left corner of Fig. 8 appears to correlate well with the strength of the CaT.

While the CaT is sensitive to  $[Ca/H]$  (Conroy & van Dokkum 2012a), its strength seems to be more closely linked to the overall metallicity  $[Z/H]$  (Brodie et al. 2012). As noted before, the chemical abundance pattern of the Vazdekis et al. (2012) stellar library at high metallicities is that of the stellar neighbourhood, while at lower metallicities many of the stars in the stellar library are GC members. In Usher et al. (2012), the Vazdekis et al. (2010) version of the Vazdekis et al. (2003) models were used to derive a relationship between CaT strength and metallicity. There, better agreement was found between the CaT-based metallicities and Lick-index-based metallicities (derived using the models of Thomas, Maraston & Bender 2003) if the model metallicities were corrected for the differences between  $\alpha$  element enhancement in the solar neighbourhood and the expected  $\alpha$  element enhancement of GCs ( $[\alpha/Fe] = 0.3$ ). After applying the same corrections to the model metallicities as Usher et al. (2012, their equation 9), we found the following relation between CaT and metallicity using the Vazdekis et al. (2012) models with an age of 12.6 Gyr and a Kroupa (2001) IMF:

$$[Z/H] = (0.431 \pm 0.011) \times \text{CaT} + (-3.604 \pm 0.072). \quad (5)$$

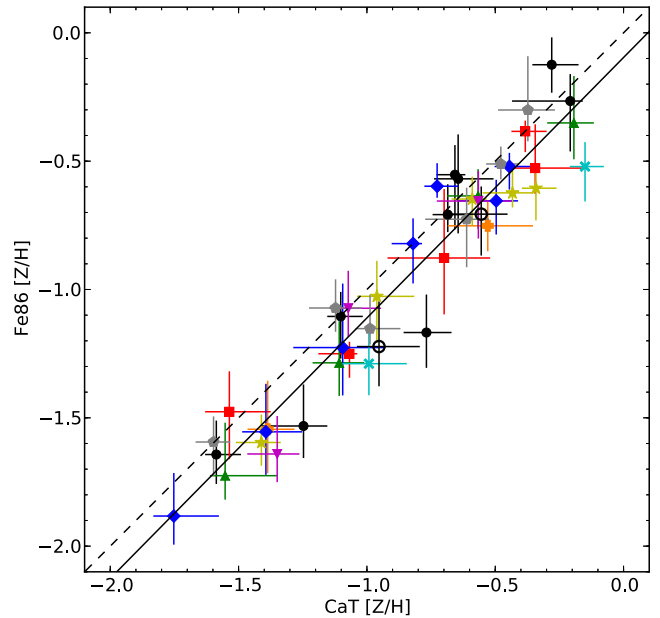
The rms of this relation is 0.027 dex.

Since the Fe86 index is dominated by a mixture of Fe and Ti lines, it is unclear whether and how metallicities derived from this index should be corrected for the abundance pattern of the models. Assuming that the index is more strongly sensitive to  $[Fe/H]$  than  $[Z/H]$ , we do not adjust the model metallicity. We derived the following relation between the strength of the Fe86 index and metallicity using the Vazdekis et al. (2012) models with an age of 12.6 Gyr and a Kroupa (2001) IMF:

$$[Z/H] = (0.604 \pm 0.012) \times \text{Fe86} + (-2.663 \pm 0.041). \quad (6)$$

The rms of this relation is 0.029 dex.

In Fig. 10, the CaT and Fe86 metallicities are compared. Although the slope of the relation between the two metallicities is consistent with unity ( $1.00 \pm 0.04$ ), the Fe86 metallicities are offset to lower metallicities by  $0.11 \pm 0.04$  dex. As mentioned above, the models of Conroy & van Dokkum (2012a) also underpredict the observed strength of the CaT by a similar amount. The rms of this fit is 0.13 dex and the  $\chi^2$  value is 34.1 for 43 degrees of freedom. Although the CaT index measures the strength of the CaT feature only, the wavelength range that the stellar templates are fitted to as part of the CaT measurement process includes most of the lines in the Fe86 index as well as the Mg88 index so the CaT index is not completely independent of either the Fe86 index or the Mg88 index. It is reassuring that spectral features dominated by different species (Ca II for the CaT, Fe I and Ti I for Fe86) give consistent results modulo a 0.1 dex offset. This suggests that both are measuring metallicity accurately. It would be desirable to use weak Fe and



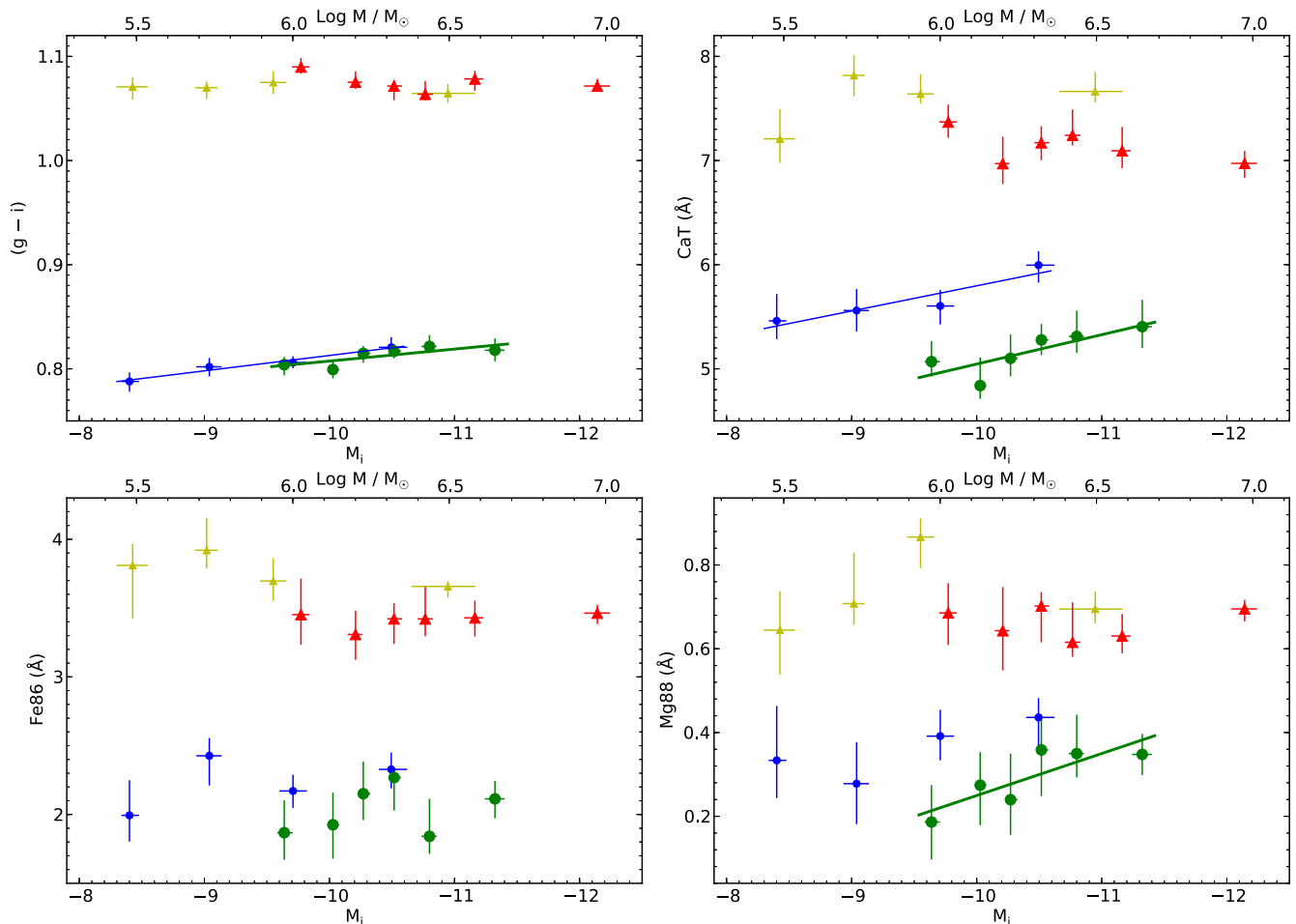
**Figure 10.** Fe86-based metallicities versus CaT based metallicities. For both indices, the 12.6 Gyr old, Kroupa (2001) IMF Vazdekis et al. (2012) model was used to convert the index strengths into metallicities. For the CaT-based metallicity, equation 9 of Usher et al. (2012) was used to correct the model metallicities for the effects of the  $[\alpha/Fe]$  pattern of the solar neighbourhood. The points are colour coded by their host galaxy using the same colours and shapes as in Fig. 8. The dashed line is 1 to 1. The solid line is a fit to the two metallicity determinations. The slope of the fit is consistent with unity ( $1.00 \pm 0.04$ ) although the intercept ( $-0.11 \pm 0.04$ ) is significantly offset from zero. The rms of the fit is 0.13 dex; the  $\chi^2$  value is 34.1 for 43 degrees of freedom.

other metal lines from across the DEIMOS wavelength range to measure metallicities of individual GCs as Kirby et al. (2008) has done for individual RGB stars.

#### 4.3 Stellar population variations between galaxies

As in Usher et al. (2012), we see variations in the  $(g-i)$ -CaT relation between galaxies (Figs 1 and 2). At blue colours the GCs of NGC 1407 and NGC 4278 have weaker CaT values than predicted by equation (1). For red colours, NGC 1407, NGC 4278, NGC 4365 and NGC 4649 GCs have weaker CaT strengths than predicted from their colours while NGC 3115, NGC 3377, NGC 4473 and NGC 4494 have stronger than predicted CaT values. The shapes of the colour-CaT relations of NGC 1407 and NGC 4494 are radically different as first noticed by Foster et al. (2011).

As seen in Fig. 8, at low CaT values NGC 3115, NGC 3377 and NGC 4473 show bluer colours compared to the other galaxies while at high CaT values NGC 3377 and NGC 4494 also show bluer colours than the other galaxies. Additionally, at low CaT values NGC 3115, NGC 3377 and NGC 4473 show higher Na82 strengths than the other galaxies. At high CaT values, NGC 3377 and NGC 4494 show slightly weaker Na82, Fe86, Mg88 and TiO89 strengths than those of the other galaxies and stronger  $H\alpha$  strengths. In the Vazdekis et al. (2012) models at high metallicities,  $(g-i)$  becomes bluer, Na82, Fe86 and Mg88 index strengths become weaker and  $H\alpha$  becomes stronger at younger ages while the CaT strength is unchanged. Since an extremely blue horizontal branch produces bluer colours, stronger  $H\beta$  indices and weaker metal index strengths in the Lick index spectral regions (Chung et al. 2013b), it is also

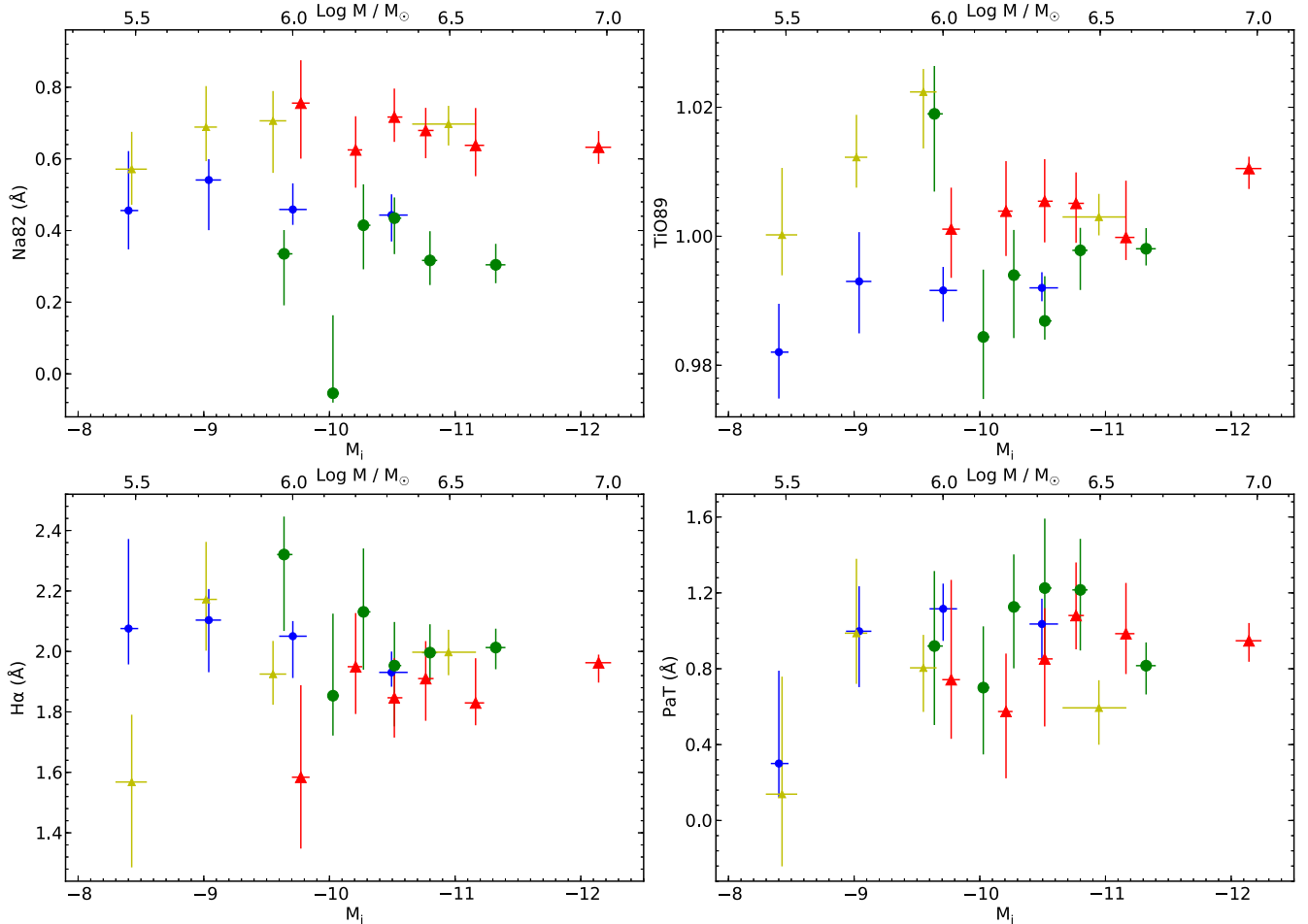


**Figure 11.** Stellar population properties of GCs stacked by magnitude in fixed colour ranges. The stacks of blue GC spectra in the faint galaxy group are plotted as small blue circles, blue GC spectra in the bright galaxy group as large green circles, red GC spectra in the faint group as yellow small triangles and red GC spectra in the bright group as red large triangles. The colours and symbols are the same as in Fig. 6. For illustrative purposes, the mass corresponding to the absolute magnitude and a consistent mass-to-light ratio of 2 is shown on the top of each plot. Where a significant relation is observed with magnitude, the relation is plotted with a line (thicker for the bright group). Top left:  $(g - i)$  colour as a function of  $i$ -band absolute magnitude. For both red and blue, the two groups have consistent colours at fixed magnitude. The blue stacks from both groups become significantly redder at brighter luminosities (the blue tilt). Top right: the CaT index as a function of  $i$ -band absolute magnitude. Both red and blue stacks show significant CaT offsets between the faint and bright groups. Both the blue faint group and bright group stacks have significantly stronger CaT values at brighter magnitudes than at fainter magnitudes. Bottom left: the Fe86 weak iron line index as a function of  $i$ -band absolute magnitude. Like the CaT index, both red and blue stacks show significant Fe86 offsets between the faint group and bright group. Bottom right: the Mg88 magnesium index as a function of  $i$ -band absolute magnitude. Like the CaT and Fe86 indices, both red and blue stacks show significant Mg88 offsets between the faint group and bright group. The Mg88 values of the bright blue stacks become significantly stronger with brighter magnitudes.

possible that blue horizontal branches are responsible. We note that while the high CaT index stack of NGC 4494 is an outlier in the colour–CaT relation, it is consistent with other galaxies in the colour–Fe86 relation (Fig. 7). Without knowing whether the CaT or Fe86 measurement better reflects the metallicity of NGC 4494’s metal-rich spectral stack, we cannot tell whether it is truly outlier.

Although their colours at fixed magnitude are consistent, in the spectra stacked by magnitude (Fig. 11), the CaT and Fe86 indices of the faint group are all significantly stronger ( $p > 0.95$ , see Table B1) at  $M_i = -10$  than those of the bright group for both the blue and red colours. If anything, the red GCs in the faint group are bluer than those in the bright group although this difference is not significant ( $p = 0.936$ ). Using equation (1), this small colour difference of  $(g - i) = -0.0101^{+0.0072}_{-0.0071}$  would correspond to a small CaT difference of  $-0.049^{+0.035}_{-0.035}$  Å which is in the opposite direction to and much smaller in magnitude than the observed CaT difference of

$0.416^{+0.138}_{-0.136}$  Å. The Mg88 indices of the faint group are also stronger than those of the bright group although the difference for the red GCs barely falls below our significance limit ( $p = 0.946$ ). There is also evidence that the faint group has stronger Na82 indices than the bright group stacks only at blue colours (Fig. 12). NGC 1407’s radial velocity places its Na82 index in the same observed wavelength as the strongest telluric absorption lines potentially biasing Na82 measurements. We restacked the bright group without NGC 1407 and measured consistent Na82 values to those of stacks including this galaxy. There is no evidence that there is a Na82 index difference between the groups for red colours. Nor is there evidence that the TiO89, H $\alpha$  or PaT index strengths are significantly different between the faint and bright groups at either blue or red colours. Since the metal index offsets between the two galaxy groups are seen at all luminosities, the offsets between the groups are not due to the index–metallicity relations changing with magnitude.



**Figure 12.** GCs stacked by magnitude in fixed colour ranges. The stacks of blue GC spectra in the faint group are plotted as small blue circles, blue GC spectra in the bright group as large green circles, red GC spectra in the faint group as yellow small triangles and red GC spectra in the bright group as red large triangles. For illustrative purposes, the mass corresponding to the absolute magnitude and a constant mass-to-light ratio of 2 is shown on the top of each plot. Top left: the Na82 sodium index as a function of  $i$ -band absolute magnitude. There is weak evidence that the faint group blue stacks have stronger Na82 index strengths than the bright group blue stacks. Top right: the TiO89 TiO index as a function of  $i$ -band absolute magnitude. Bottom left: the  $H\alpha$  index as a function of  $i$ -band absolute magnitude. Bottom right: the PaT Paschen line index as a function of  $i$ -band absolute magnitude. None of the Na82, TiO89,  $H\alpha$  or PaT indices have any significant index–magnitude relationships.

Since the strengths of Fe86 and Mg88 correlate strongly with the strength of the CaT, we checked to see if the Fe86 and Mg88 offsets at  $M_i = -10$  are consistent with being caused by the CaT offsets at the same magnitude ( $0.75_{-0.18}^{+0.10}$  and  $0.42_{-0.14}^{+0.14}$  Å for the blue and red stacks, respectively). Using equation (3), the CaT offsets predict Fe86 offsets of  $0.54_{-0.10}^{+0.07}$  and  $0.30_{-0.10}^{+0.10}$  Å, consistent with the observed values of  $0.31_{-0.15}^{+0.14}$  and  $0.33_{-0.15}^{+0.11}$  Å. Likewise, using equation (4), the CaT offsets predict Mg88 offsets of  $0.138_{-0.032}^{+0.018}$  and  $0.077_{-0.025}^{+0.025}$  Å, again consistent with observed values of  $0.146_{-0.058}^{+0.056}$  and  $0.077_{-0.050}^{+0.065}$  Å, respectively. The consistency of these three metal indices and the strong correlation of the CaT and Fe86 indices with metallicity points to a metallicity difference between the two groups. It also suggests that there is no large variation in the  $[\text{Mg}/\text{Fe}]$ – $[\text{Fe}/\text{H}]$  or  $[\text{Ca}/\text{Fe}]$ – $[\text{Fe}/\text{H}]$  relations between galaxies. At low CaT strengths, the top-right panel of Fig. 9 shows that the TiO89 strength is largely independent of CaT strength in line with the lack of a significant difference in TiO89 between the two groups. However, at CaT strengths above  $\sim 6.3$  Å, the TiO89 strength increases with CaT strength. Fitting the slope of the CaT–TiO89 relation above CaT = 6.3 Å to be  $0.0166 \pm 0.0048$  Å $^{-1}$ , the CaT difference

predicts a TiO89 difference of  $0.0069_{-0.0023}^{+0.0023}$  which is consistent with the observed, non-significant ( $p = 0.939$ ), TiO89 difference of  $0.0073_{-0.0045}^{+0.0043}$ . The behaviour of the TiO89 index, which is sensitive to the O and Ti abundances, provides further evidence that the observed metal index differences are caused by variations in the colour–metallicity relation. Since the  $H\alpha$  and PaT indices show weak or no dependence on CaT at low CaT values, the lack of offsets between the faint and bright groups for these indices is consistent with the CaT, Fe86 and Mg88 offsets being due to a metallicity offset.

The behaviour of the Na82 index is difficult to interpret with existing models and data. Besides the Na I doublet at 8183 and 8195 Å, the index definition contains a TiO molecular bandhead and is adjacent to the Paschen limit at 8205 Å. At solar metallicity the index is sensitive to age, the IMF and  $[\text{Na}/\text{Fe}]$  (Conroy & van Dokkum 2012a). As mentioned above, the Vazdekis et al. (2012) models are unreliable in the Na82 spectral region at low metallicities. Besides the model uncertainties, the Na82 index likely suffers from measurement systematics due to uncertain telluric corrections. As seen in the left panel of the middle row of Fig. 8, at low CaT strengths,

there is no significant relation (a slope of  $0.02 \pm 0.07$ ) between the CaT and Na82 indices while above CaT = 6.3 Å a relation is seen with a slope of  $0.29 \pm 0.12$ . These CaT–Na82 relations predict Na82 differences of  $0.016_{-0.004}^{+0.002}$  and  $0.121_{-0.039}^{+0.040}$  for blue and red colours, respectively, yet differences of  $0.201_{-0.074}^{+0.063}$  and  $-0.019_{-0.075}^{+0.079}$  are observed. The larger than expected difference in Na82 between galaxy groups for blue GCs could be explained by higher Na abundances in the faint group while for red GCs younger ages in the faint group could explain the smaller than expected difference. The behaviour of the Na82 index does not seem to be following the behaviour of the other metal indices.

Although the spectral indices have been measured in a homogeneous manner using the same spectroscopic setup, data reduction pipeline and analysis code, the photometry is heterogeneous with data from different telescopes, analysed in different ways. If the differences between galaxies were caused by incorrect photometric zero-points or by incorrectly removing the foreground extinction, then the observed colours would be shifted bluer or redder. However, the differences in colour–CaT relationships cannot be simply resolved by shifting individual galaxies’ GC colours. If a colour term has been neglected in the zero-points, the slope of the colour–CaT relationship could change. While this could improve the agreement for NGC 3377 and NGC 4494, changing the slope would not help NGC 1407. It is unlikely that reddening would be correlated with metallicity nor can extinction make an object bluer. For NGC 3115 (Jennings et al. 2014), NGC 4278 (Usher et al. 2012), NGC 4365 (Blom et al. 2012), NGC 4473 (Jordán et al. 2009) and NGC 4649 (Strader et al. 2012) where ACS  $g_z$  photometry is available, the same colour–CaT relations are seen for both the ACS and Suprime-Cam photometry. Furthermore, photometric issues cannot explain why CaT weak spectra from NGC 3115, NGC 3377 and NGC 4473 have stronger Na82 strengths than other galaxies and why CaT strong NGC 3377 and NGC 4494 spectra have weaker metal and stronger  $H\alpha$  indices than other galaxies. Although it is highly desirable to improve the uniformity of the SLUGGS GC photometry, we do not believe heterogeneous photometry to be the source of the colour–metal index variations.

An important caveat that the median magnitude of GCs in this study is  $M_i = -10.3$  which is significantly brighter than the turnover magnitude ( $M_i = -8.4$ ; Villegas et al. 2010) so that stellar populations of the observed sample of GCs may not be representative of all GCs in the sample galaxies. Since in several galaxies including NGC 1407 (Romanowsky et al. 2009), M87 (Strader et al. 2011b) and NGC 5128 (Woodley et al. 2010a), GCs with magnitudes comparable to or brighter than  $\omega$  Cen ( $M_i = -11.3$ ; Vanderbeke et al. 2014a) show different kinematics than fainter GCs, it is quite possible that the very brightest GCs are a distinct population with different stellar population trends than the majority of GCs. In this study, 9 per cent of the GCs are brighter than  $\omega$  Cen.

Although the mean magnitude of GCs in the galaxies with bluer, more linear colour–CaT relations is fainter, we do not believe that the differences in colour–CaT relations between galaxies are due to stellar population variations with GC magnitude. If we restrict the CaT measurements of individual GCs in Fig. 1 to GCs in the magnitude range  $-10.5 < M_i < -9.5$ , we see the same colour–CaT relationships. In Fig. 11, stacked spectra with the same colours and absolute magnitudes from the two galaxy groups have significantly different metal index strengths. This is also seen in the two groups having consistent colour– and index–magnitude slopes. Together these suggest that the differences in the luminosities of observed GCs between galaxies are not responsible for the observed colour–CaT variations.

With the small number of galaxies in this study, it is difficult to reliably study how the colour–metal index variations correlate with host galaxy properties. The three most luminous galaxies in this study (NGC 1407, NGC 4365 and NGC 4649) belong to the bright group while least luminous galaxy in our sample (NGC 3377) belongs to the faint group. The remaining four galaxies all have similar luminosities ( $M_K \sim -24$ ) with only NGC 4278 belonging to the bright group. This hints, but does not prove, that the shape of the colour–CaT relation is linked to galaxy mass or a property that depends on galaxy mass.

Possible stellar population differences that could explain the observed variations in metal index strength at fixed colour between galaxies include different ages, chemical abundance patterns or IMFs. Before discussing each of these in turn, it should be noted that one way that the age and abundance pattern of a stellar population manifest themselves is through the morphology of its horizontal branch. In Milky Way GCs, metallicity, age and helium abundance seem to determine the horizontal branch morphology (Gratton et al. 2010; Milone et al. 2014). It should also be noted that our results only indicate differences in the mean index–colour relations from galaxy to galaxy and that there likely exists spreads in some of these properties within each galaxy (e.g. a range of GC ages at fixed metallicity as is seen in Milky Way GCs; Dotter, Sarajedini & Anderson 2011; VandenBerg et al. 2013).

#### 4.3.1 Age differences?

One possible explanation for the colour–metal index variations is differences in the age–metallicity relation between galaxies. As can be seen in the top-left panel of Fig. 8, younger ages lead to bluer colours. The Vazdekis et al. (2012) models predict that age has little effect on the CaT at high metallicities but at low metallicities predict stronger CaT values at younger ages. Both the Conroy et al. (2009) and Chung et al. (2013a) models predict that for solar iron abundances 8 Gyr ( $z \sim 1$ ) populations are 0.1 mag bluer in ( $g - i$ ) than 13 Gyr ( $z \sim 8$ ) populations. At low metallicities, the morphology of horizontal branch strongly depends on age (e.g. Dotter et al. 2010). The horizontal branch morphology has a large effect on the integrated colours while having a relatively minor effect on CaT strengths (Maraston 2005; Conroy et al. 2009; Brodie et al. 2012). For example, the Chung et al. (2013a) models, which include a detailed treatment of the horizontal branch, predict that 12 Gyr ( $z \sim 4$ ) populations are 0.05 mag bluer in ( $g - i$ ) than 13 Gyr ( $z \sim 8$ ) populations at  $[\text{Fe}/\text{H}] = -2$ .

The bluer colours of the GCs in the faint group of galaxies suggest that less massive galaxies might have younger GCs on average. The concept of GC starting to form slightly later and continuing to form over a period of time is consistent with observations of single stellar population equivalent ages being younger in less massive galaxies (e.g. Thomas et al. 2005; Kuntschner et al. 2010; Smith et al. 2012). Forbes et al. (2007) compared galaxy ages with the mean colours of their GC subpopulations. After removing the known relationships between GC colour and host galaxy luminosity, they found no relationships between the galaxy central age and GC colour. They interpret this as either GCs in younger galaxies are more metal rich than those in old galaxies or that GCs are uniformly old. However, by removing the GC colour–galaxy luminosity relations to subtract the GC metallicity–galaxy luminosity relations, they also take away any GC age–galaxy luminosity relations that might be present.

Spectroscopic studies (e.g. Cohen, Blakeslee & Ryzhov 1998; Forbes et al. 2001; Puzia et al. 2005; Strader et al. 2005; Norris

et al. 2008) of GC ages have shown that the majority of GCs have old ( $\sim 10$  Gyr) ages and that there is no significant difference in age between the metal-poor and metal-rich populations. Evidence for a population of young or intermediate-age metal-rich GCs have been observed in galaxies such as NGC 5128 (e.g. Woodley et al. 2010b) and NGC 4649 (Pierce et al. 2006). In one of the only studies of the stellar populations of GCs in a sub- $L^*$  galaxy (NGC 7457,  $M_K = -22.4$ ), Chomiuk, Strader & Brodie (2008) found that the two most metal-rich GCs they observed display emission lines which likely indicates young ages. In their study of GCs in Local Group dwarf ellipticals, Sharina, Afanasiev & Puzia (2006) derived significantly younger ages for metal-rich GCs compared to metal-poor GCs in the same galaxies. A caveat to all these studies is the difficulty in determining accurate ages for old stellar populations from integrated spectra, especially due to the effects of the horizontal branch. Additionally, the Milky Way GCs thought to have formed in dwarf galaxies show young ages at fixed metallicity than GCs thought to have formed *in situ* (Forbes & Bridges 2010; Leaman, VandenBerg & Mendel 2013). Taken together, this suggests that it is plausible that less massive galaxies have slightly younger GCs systems which would be bluer.

We can use the CaT index strengths and  $(g - i)$  colours of the two galaxy groups to estimate the mean age differences required to explain the differences in CaT index observed between the groups. Using equation (5) and the Chung et al. (2013a) models, we estimate ages of 7.3 and 13.1 Gyr and metallicities of  $[Z/H] = -1.1$  and  $-1.4$  for the blue GCs in the faint and bright galaxy groups, respectively. For the red GCs, we estimate 7.8 and 10.6 Gyr and metallicities of  $[Z/H] = -0.3$  and  $-0.5$  for the faint and bright groups, respectively. Using the Conroy et al. (2009) models, we get ages of 6.6 and 8.3 Gyr for the blue colour range, and of 6.5 and 12.0 Gyr for the red colour range. These age and metallicities assume that the relationship between CaT strength and metallicity does not vary with age and that the observed differences in the colour–CaT relationship are solely due to age. The uncertainties inherent in stellar population modelling are highlighted by the differences in ages derived using the Chung et al. (2013a) and Conroy et al. (2009) models. It should be noted that since the S/N cut we apply to the spectra effectively imposes a magnitude limit and since younger GCs are brighter than old GCs due to the effects of stellar and dynamical evolution, our sample is biased towards younger GCs and the true difference in mean ages may be smaller.

As less luminous galaxies have on average younger stellar populations, it is quite possible that they host on average younger GCs. Under the assumption that the differences in CaT strength are solely due to an age difference, an age difference of only a few Gyr is required between the two galaxy groups. Differences in the GC age–metallicity relation between galaxies are a possible explanation for at least some of the variation in the colour–metal index relation.

#### 4.3.2 Abundance differences?

The abundance pattern of a stellar population also affects its colours. Helium enhanced populations have hotter isochrones and horizontal branches which produce bluer optical colours. In models of Chung et al. (2013b), a He enhanced population ( $Y = 0.33$ ) is  $\sim 0.25$  mag bluer in  $(g - z)$  than a He normal population ( $Y = 0.23$ ) at solar iron abundance but no significant difference in colour is seen below  $[Fe/H] = -1$ . The effect of  $\alpha$ -element enhancement depends greatly on the colour considered. In the models of Coelho et al. (2007),  $\alpha$ -element enhanced populations are bluer than scaled solar

populations for the bluest optical colours (e.g.  $\Delta(u - g) = -0.04$  for  $[Fe/H] = 0$  and an age of 12 Gyr) but are redder than scaled solar populations at fixed  $[Fe/H]$  for the redder colours (e.g.  $\Delta(r - z) = 0.07$ ). Different elements also affect optical colours in different ways. Lee et al. (2009) found that increased abundances of heavier elements such as Mg, Ti and Fe leads to redder colours while enhanced C, N or O leads to bluer colours.

While we do not believe that  $[Mg/Fe]$  or  $[Ca/Fe]$  varies significantly between galaxies due to the lack of differences in the Mg88–CaT or Fe86–CaT relations, there still could be variations in the abundances of He, C, N or O between the two groups. In the Milky Way, GCs contain stars with both ‘normal’  $\alpha$ -element enhanced abundance patterns and stars that are relatively He, N and Na rich and C and O poor. It is possible that the process that produces these abundance anticorrelations could vary from galaxy to galaxy, producing variations in the mean abundances of these light elements. Goudfrooij & Kruijssen (2013) found that a solar metallicity population with a normal  $\alpha$ -element enhanced abundance pattern was 0.05 mag redder in  $(g - i)$  than a He, N and Na enhanced and C and O depleted population. However, they did not consider the effects of the abundance variations on the horizontal branch. Although CaT strengths are relatively insensitive to the horizontal branch morphology, the effects of helium on the isochrones or the effects of CNO abundances on the pseudo-continuum by CN and TiO molecular bands could affect the CaT strengths. In the left panel of the middle row of Fig. 8, the Na82 index is stronger in galaxies in the faint group compared to the bright group at low CaT values. This suggests that the Na abundance is higher in the faint group than in the bright group. Variations in the abundances of elements such as He, C, N and O are capable of producing colour variations of  $\Delta(g - i) \sim 0.1$  and could be responsible for colour–metal index variations observed.

#### 4.3.3 IMF differences?

As can be seen in the upper-left panel of Fig. 8, populations rich in low-mass stars have redder colours and weaker CaT strengths. If, like the IMF of their host galaxies, the IMF of GCs was more dwarf rich in more massive galaxies then more massive galaxies would have GCs with redder colours and weaker CaT values compared to less massive galaxies in agreement with the colour–CaT variations seen. However from the combination of CaT, Na82 and Fe86 indices, we do not see evidence that any of the galaxies host low-mass star rich IMFs. Due to the effects of internal dynamical evolution, GCs preferentially lose low-mass stars (e.g. Hénon 1969; Baumgardt & Makino 2003; Kruijssen 2009). Thus, their present-day stellar mass distributions are different from those of idealized single stellar populations with the same IMF. Removing low-mass stars would make a GC bluer and appear to have a bottom light IMF which would translate into a stronger CaT value. However, as can be seen in the upper-left panel of Fig. 8, the removal of low-mass stars does not have a large enough effect on colour or CaT strength to explain the observed galaxy to galaxy variations. We consider it unlikely that variations in the IMF or the present-day mass function between galaxies can explain the observed metal line index–colour variations.

#### 4.3.4 Summary

From the colour–CaT relations of individual GCs and from the differences in metal index strengths of stacked spectra with the same colours and luminosities, there is strong evidence that the



GC colour–metallicity relation varies from galaxy to galaxy. The data hint that lower mass galaxies have bluer, more linear colour–metallicity relations than higher mass galaxies. We do not believe that these variations are due to heterogeneous photometry or differences in GC luminosities between galaxies. Homogeneous colour and metallicity data from more galaxies are required to confirm whether the variations in the colour–metallicity relation are real and to establish whether they are linked to galaxy mass or a related galaxy property such as stellar age. We consider a variation in GC ages between galaxies to be a possible explanation for the observed colour–metallicity relations. Alternatively, variations in the abundances of He, C, N or O between galaxies could cause the colour–metal line index relations to vary. These two scenarios are not exclusive. We do not consider differences in the IMF to be a likely cause.

It would be highly desirable to repeat the spectral stacking procedure with spectra at blue to try to understand if the colour–metal index relation variations are caused by GC age variations. Blue spectra would also allow us to look for variations in C or N abundance by studying the strengths of the associated molecular bands. Acquiring photometry over a wider wavelength range would also be useful in understanding the origin of the colour–metal index variations. For example, optical–near-infrared colours such as  $(i - K)$  are less sensitive to the effects of age and the horizontal branch morphology compared to  $(g - i)$  while near-ultraviolet colours such as  $(u - g)$  are more sensitive (e.g. Conroy et al. 2009). Since abundance variations in MW GC stars are most apparent in the bluest filters (e.g. Monelli et al. 2013),  $u$ -band photometry would also be useful to distinguish abundance variations from age variations. For example, the calculations of Goudfrooij & Kruijssen (2013) predict a He, N and Na rich, C and O poor population to be 0.05 mag redder in  $(u - g)$  and 0.05 mag bluer in  $(g - i)$  than a population with a normal  $\alpha$ -enhanced abundance pattern. In the Conroy et al. (2009) models, a 9 Gyr solar metallicity population is 0.10 and 0.05 mag bluer in  $(u - g)$  and  $(g - i)$ , respectively, than a 12 Gyr population.

#### 4.4 Stellar population variation with luminosity

As seen in Fig. 11, both the bright and faint group blue stacks are significantly redder at brighter magnitudes. This is the blue tilt that is seen in Fig. 5. The slopes seen in the two groups are consistent with one another. The observed blue tilt is smaller but consistent with the mean blue tilt observed by Mieske et al. (2010) despite the narrow colour range used for stacking. The red stacks show no evidence for a colour–magnitude relation which is consistent with its non-detection in photometric studies such as Harris (2009), although Mieske et al. (2010) do observe a weak red tilt in the centres of their most massive galaxies.

Both the blue faint group and the blue bright group CaT values become significantly stronger with magnitude. Although the relations are offset from one another, the slopes are consistent. No trend with magnitude is seen for the red spectra stacks. Although no significant Fe86–magnitude trends are observed, both the bright and faint group blue stacks show slightly higher Fe86 at brighter magnitudes. Both the blue faint groups and bright groups show higher Mg88 values at brighter magnitudes although only the bright groups shows a significant relation. The red stacks show no Mg88–magnitude relation. As seen in Fig. 12, none of the Na82, Ti89, H $\alpha$  or PaT indices show any significant relations with magnitude.

Using equation (1) and the observed colour–magnitude relations, we predict CaT–magnitude relation slopes of  $-0.148^{+0.051}_{-0.062}$  and

$-0.115^{+0.064}_{-0.086}$  for the faint and bright groups, respectively. These predicted slopes are consistent with the observed CaT–magnitude slopes. Using the equation (3) and the observed CaT–magnitude relation, we predict the Fe86–magnitude slopes to be  $-0.175^{+0.099}_{-0.061}$  and  $-0.204^{+0.125}_{-0.101}$  for the faint and bright groups, respectively, consistent with the observed Fe86–magnitude slopes. Likewise, using equation (4), the observed CaT–magnitude relations predict Mg88–magnitude slopes of  $-0.045^{+0.025}_{-0.015}$  and  $-0.052^{+0.032}_{-0.026}$ , again consistent with the observed relations. Since at low CaT values we do not see any significant relation between CaT index strength and any of the Na82, Ti89, H $\alpha$  or PaT indices in the spectra stacked by colour, the lack of significant trends with magnitude for these indices is consistent with the CaT–magnitude trends we observe. As the CaT index strongly correlates with metallicity, the colour– and index–magnitude relations are all consistent with a metallicity–magnitude relation.

Since the slopes of the metal line index–magnitude relations are consistent with being explained by the observed colour–metallicity relations, we do not see any evidence for a changing colour–metallicity relation with magnitude. The colour– and index–magnitude slopes are consistent between the faint and bright groups, indicating that the blue tilt does not greatly vary from galaxy to galaxy. However, photometric studies (e.g. Mieske et al. 2010) have found that the blue tilt is stronger in centres of giant galaxies than in the outskirts of giant galaxies or in dwarf galaxies. The GCs in this study are on average further away from centres of their host galaxies than those in Mieske et al. (2010) so it is not surprising that our colour–magnitude slopes are closer to the those measured by Mieske et al. (2010) in the outer regions of giant galaxies. For the same reason, the lack of a red tilt in our data is also consistent with Mieske et al. (2010) as they only find evidence for a red tilt in the centres of massive galaxies.

The observed trends are all consistent with the blue tilt being purely a metallicity–mass trend with no variations in age, element enhancement or IMF with magnitude. However, the uncertainties in our trends with magnitude are large so weak trends could still exist. The lack of a bottom heavy IMF in the most luminous GCs (Section 4.1) and the consistency of the colour–magnitude trend with being a metallicity–magnitude trend suggest that internal dynamical evolution does not play a major role in producing the observed colour–magnitude trends contrary to the suggestions by Goudfrooij & Kruijssen (2014).

This is not the first time the blue tilt has been observed spectroscopically; Foster et al. (2010) found that brighter blue GCs in NGC 1407 have stronger CaT values than fainter blue GCs. We see the same CaT trend extending to 2 mag fainter along with consistent trends in other metal indices. Schiavon et al. (2013) studied the variation of line indices and abundances with magnitude in metal-rich GCs in M31. They saw no variation in [Fe/H], [Mg/Fe] or [C/Fe] with mass but saw higher [N/Fe] abundances and stronger NaD indices at higher masses. We note that GCs in our sample are brighter on average than those in the Schiavon et al. (2013) study.

Making the assumption that Mg88 follows [Mg/H], that the Mg88–magnitude relation is consistent with the CaT–magnitude relation suggests that [Mg/Fe] does not vary significantly with magnitude, consistent with the results of Schiavon et al. (2013). As mention above, the Na82 index is difficult to interpret. Under the assumption of no IMF or age variation with magnitude, the lack of a Na82–magnitude gradient could be seen as evidence of a lack of any [Na/Fe]–magnitude relationship. However, the NaD–mass relation observed by Schiavon et al. (2013) likely corresponds to

a Na82–magnitude relation weak enough to be consistent with our observations. If the blue tilt is caused by an extension of the light element variations seen in Milky Way GCs to higher masses, then we would expect light element enhancements to vary with luminosity. The [N/Fe] and NaD trends observed by Schiavon et al. (2013) at high metallicities are consistent with this picture. Repeating the our stacking procedure with blue optical spectra would allow us to measure how [C/Fe] and [N/Fe] vary with magnitude by studying their molecular bands.

## 5 CONCLUSIONS

Here, we have stacked 1137 GC spectra from 10 galaxies to understand their stellar populations. Within each galaxy we stacked spectra by GC colour. To disentangle the effects of the varying GC colour–CaT relation from trends with magnitude, we combined spectra from the four galaxies (NGC 3115, NGC 3377, NGC 4473 and NGC 4494) with bluer colour–CaT relations together in one group (which we name the faint group) and spectra from the four galaxies with redder colour–CaT relations (NGC 1407, NGC 4278, NGC 4365 and NGC 4649) in another group (the bright group). This was done before stacking red and blue spectra from each group by magnitude. For each stacked spectrum, we calculated the mean colour and magnitude and measured the velocity dispersion and strengths of 7 spectral indices. Since the colour and CaT values of the spectra stacked by colour agree with the mean values of individual spectra (Fig. 1) and the velocity dispersions of the spectra stacked by magnitude agree with the literature values (Fig. 6), we are confident of the stacking procedure.

Comparing our measurements with the simple stellar population models of Vazdekis et al. (2012), we find that the models generally agree with the observed relations between the spectral indices and ( $g - i$ ) colour. However, the models do not predict the shapes of the observed colour–CaT or TiO89–CaT relations and predict stronger H $\alpha$  and PaT indices than observed. Using the colours plus the CaT, Na82, Fe86, H $\alpha$  and PaT strengths of the stacked spectra, we find the observed GCs to be consistent with old ages and a bottom light IMF. Using the strengths of weak metal lines in the CaT region (i.e. the Fe86 index) to measure metallicity shows promise (Fig. 10, Section 4.2). However, better stellar population models at low metallicity which include the ability to vary chemical abundances and the horizontal branch morphology are needed to fully utilize our data.

In line with the previous studies by Foster et al. (2011) and Usher et al. (2012), in Fig. 1 large variations in the colour–CaT relation can be seen. Despite having consistent colours, the CaT, Fe86 and Mg88 indices are stronger in stacked spectra from the faint group compared to the bright group stacked spectra at both red and blue colours at fixed luminosity. The sizes of these offsets are consistent with a metallicity difference between the two groups. The combination of the variation in colour–CaT index relations and the metal index offsets seen between the groups of galaxies at fixed colours and luminosities makes it clear that the colour–metallicity relation varies from galaxy to galaxy. Variations in GC ages might explain the observed colour–metallicity variations. Interestingly, the galaxies with bluer colour–metallicity relations have on average fainter luminosities. This could indicate that GCs in less massive galaxies started forming slightly later and continued to form a few Gyr longer than in more massive galaxies. Alternatively, differences in the average abundances of light elements such as He, C, N and O could produce the variations in colour–metallicity relations observed. To

understand if age or abundance variations are the explanation for the observed variations between optical colour and metal index strengths, a similar stacking study with spectra at blue wavelengths would be useful as would ultraviolet and near-infrared photometry.

As seen in Fig. 11, at low metallicities we see that GC colours become redder and the CaT, Fe86 and Mg88 metal indices become stronger with brighter magnitudes. The colour– and metal index–magnitude relations are all consistent with the blue tilt observed in photometric studies such as Strader et al. (2006), Harris (2009) and Mieske et al. (2010). The blue tilt appears to be just a mass–metallicity relationship with no major change in the IMF or chemical abundances with mass.

In summary, by stacking Keck DEIMOS spectra we find that most GCs are old and have Milky Way-like IMFs. We see significant differences between galaxies in the strengths of several metal indices at fixed colour and luminosity which strongly suggests that the GC colour–metallicity relation is not universal. The origin of this variation is presently unknown although age or abundance variations are possible explanations. For metal-poor GCs, we also see that GC colours become redder and metal line indices become stronger with brighter magnitudes in line with the frequently observed blue tilt.

## ACKNOWLEDGEMENTS

We thank Michael Murphy for valuable discussions as well as Gary Da Costa and John Blakeslee for their useful comments. We thank Sreeja Kartha for her assistance with the NGC 2768 photometry. We thank the referee, Guy Worthey, for his valuable comments which greatly improved the manuscript. This research was based on data from the MILES project. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration (NASA). The analysis pipeline used to reduce the DEIMOS data was developed at UC Berkeley with support from National Science Foundation (NSF) grant AST-0071048. This publication made use of PYRAF and PYFITS which are products of the Space Telescope Science Institute (STScI), which is operated by AURA for NASA. This research made use of TOPCAT (Taylor 2005). Some of the data presented herein were obtained at the W. M. Keck Observatory, operated as a scientific partnership among the California Institute of Technology, the University of California and NASA, and made possible by the generous financial support of the W. M. Keck Foundation. The authors wish to recognize and acknowledge the very significant cultural role and reverence that the summit of Mauna Kea has always had within the indigenous Hawaiian community. We are most fortunate to have the opportunity to conduct observations from this mountain. This research is based in part on data collected at Subaru Telescope, which is operated by the National Astronomical Observatory of Japan. This research is based on observations made with the NASA/ESA *Hubble Space Telescope*, obtained from the data archive at the STScI. STScI is operated by the Association of Universities for Research in Astronomy, Inc. under NASA contract NAS 5-26555. This material is based upon work supported by the Australian Research Council under grant DP130100388 and by the NSF under grants AST-1109878 and AST-1211995. Part of this work was performed on the gSTAR national facility and the swinSTAR supercomputer at Swinburne University of Technology. gSTAR is funded by Swinburne University of Technology and the Australian Government’s Education Investment Fund.

## REFERENCES

- Andrae R., 2010, preprint ([arXiv:1009.2755](https://arxiv.org/abs/1009.2755))
- Bailin J., Harris W. E., 2009, *ApJ*, 695, 1082
- Baumgardt H., Makino J., 2003, *MNRAS*, 340, 227
- Blom C., Spitler L. R., Forbes D. A., 2012, *MNRAS*, 420, 37
- Brodie J. P., Huchra J. P., 1991, *ApJ*, 379, 157
- Brodie J. P., Romanowsky A. J., Strader J., Forbes D. A., 2011, *AJ*, 142, 199
- Brodie J. P., Usher C., Conroy C., Strader J., Arnold J. A., Forbes D. A., Romanowsky A. J., 2012, *ApJ*, 759, L33
- Brodie J. P. et al., 2014, *ApJ*, in press ([arXiv:1405.2079](https://arxiv.org/abs/1405.2079))
- Cappellari M., Emsellem E., 2004, *PASP*, 116, 138
- Cappellari M. et al., 2011, *MNRAS*, 413, 813
- Cappellari M. et al., 2012, *Nature*, 484, 485
- Carretta E., 2006, *AJ*, 131, 1766
- Carretta E., Bragaglia A., Gratton R. G., Recio-Blanco A., Lucatello S., D’Orazi V., Cassisi S., 2010, *A&A*, 516, A55
- Carretta E., Lucatello S., Gratton R. G., Bragaglia A., D’Orazi V., 2011, *A&A*, 533, A69
- Cenarro A. J., Cardiel N., Gorgas J., Peletier R. F., Vazdekis A., Prada F., 2001, *MNRAS*, 326, 959
- Cenarro A. J., Cardiel N., Vazdekis A., Gorgas J., 2009, *MNRAS*, 396, 1895
- Chomiuk L., Strader J., Brodie J. P., 2008, *AJ*, 136, 234
- Chung C., Yoon S.-J., Lee S.-Y., Lee Y.-W., 2013a, *ApJS*, 204, 3
- Chung C., Lee S.-Y., Yoon S.-J., Lee Y.-W., 2013b, *ApJ*, 769, L3
- Coelho P., Bruzual G., Charlot S., Weiss A., Barbuy B., Ferguson J. W., 2007, *MNRAS*, 382, 498
- Cohen J. G., Blakeslee J. P., Ryzhov A., 1998, *ApJ*, 496, 808
- Conroy C., Spergel D. N., 2011, *ApJ*, 726, 36
- Conroy C., van Dokkum P., 2012a, *ApJ*, 747, 69
- Conroy C., van Dokkum P. G., 2012b, *ApJ*, 760, 71
- Conroy C., Gunn J. E., White M., 2009, *ApJ*, 699, 486
- Conroy C., Graves G. J., van Dokkum P. G., 2014, *ApJ*, 780, 33
- Cooper M. C., Newman J. A., Davis M., Finkbeiner D. P., Gerke B. F., 2012, *Astrophysics Source Code Library*, record ascl:1203.003
- D’Ercole A., D’Antona F., Ventura P., Vesperini E., McMillan S. L. W., 2010, *MNRAS*, 407, 854
- Da Costa G. S., Held E. V., Saviane I., 2014, *MNRAS*, 438, 3507
- Dalessandro E., Schiavon R. P., Rood R. T., Ferraro F. R., Sohn S. T., Lanzoni B., O’Connell R. W., 2012, *AJ*, 144, 126
- Decressin T., Meynet G., Charbonnel C., Prantzos N., Ekström S., 2007, *A&A*, 464, 1029
- Dotter A. et al., 2010, *ApJ*, 708, 698
- Dotter A., Sarajedini A., Anderson J., 2011, *ApJ*, 738, 74
- Faber S. M. et al., 2003, in Iye M., Moorwood A. F. M., eds, *Proc. SPIE Conf. Ser. Vol. 4841, Instrument Design and Performance for Optical/Infrared Ground-based Telescopes*. SPIE, Bellingham, p. 1657
- Forbes D. A., Bridges T., 2010, *MNRAS*, 404, 1203
- Forbes D. A., Beasley M. A., Brodie J. P., Kissler-Patig M., 2001, *ApJ*, 563, L143
- Forbes D. A., Proctor R., Strader J., Brodie J. P., 2007, *ApJ*, 659, 188
- Forbes D. A., Lasky P., Graham A. W., Spitler L., 2008, *MNRAS*, 389, 1924
- Forbes D. A., Ponman T., O’Sullivan E., 2012, *MNRAS*, 425, 66
- Foster C., Forbes D. A., Proctor R. N., Strader J., Brodie J. P., Spitler L. R., 2010, *AJ*, 139, 1566
- Foster C. et al., 2011, *MNRAS*, 415, 3393
- Freeman K. C., Rodgers A. W., 1975, *ApJ*, 201, L71
- Gallazzi A., Charlot S., Brinchmann J., White S. D. M., 2006, *MNRAS*, 370, 1106
- Girardi L., Bressan A., Bertelli G., Chiosi C., 2000, *A&AS*, 141, 371
- Goudfrooij P., Kruijssen J. M. D., 2013, *ApJ*, 762, 107
- Goudfrooij P., Kruijssen J. M. D., 2014, *ApJ*, 780, 43
- Gratton R. G., Carretta E., Bragaglia A., Lucatello S., D’Orazi V., 2010, *A&A*, 517, A81
- Gratton R. G., Carretta E., Bragaglia A., 2012, *A&AR*, 20, 50
- Harris W. E., 2009, *ApJ*, 699, 254
- Harris W. E., Whitmore B. C., Karakla D., Okoń W., Baum W. A., Hanes D. A., Kavelaars J. J., 2006, *ApJ*, 636, 90
- Henon M., 1969, *A&A*, 2, 151
- Jennings Z. G. et al., 2014, *AJ*, 148, 32
- Johansson J., Thomas D., Maraston C., 2012, *MNRAS*, 421, 1908
- Jordán A. et al., 2009, *ApJS*, 180, 54
- Kartha S. S., Forbes D. A., Spitler L. R., Romanowsky A. J., Arnold J. A., Brodie J. P., 2014, *MNRAS*, 437, 273
- Kirby E. N., Guhathakurta P., Sneden C., 2008, *ApJ*, 682, 1217
- Kirby E. N., Lanfranchi G. A., Simon J. D., Cohen J. G., Guhathakurta P., 2011a, *ApJ*, 727, 78
- Kirby E. N., Cohen J. G., Smith G. H., Majewski S. R., Sohn S. T., Guhathakurta P., 2011b, *ApJ*, 727, 79
- Kroupa P., 2001, *MNRAS*, 322, 231
- Kruijssen J. M. D., 2009, *A&A*, 507, 1409
- Kuntschner H. et al., 2010, *MNRAS*, 408, 97
- La Barbera F., Ferreras I., Vazdekis A., de la Rosa I. G., de Carvalho R. R., Trevisan M., Falcón-Barroso J., Ricciardelli E., 2013, *MNRAS*, 433, 3017
- Leaman R., VandenBerg D. A., Mendel J. T., 2013, *MNRAS*, 436, 122
- Lee H.-c. et al., 2009, *ApJ*, 694, 902
- Maraston C., 2005, *MNRAS*, 362, 799
- Marino A. F. et al., 2011a, *A&A*, 532, A8
- Marino A. F. et al., 2011b, *ApJ*, 731, 64
- Mei S. et al., 2007, *ApJ*, 655, 144
- Mieske S. et al., 2010, *ApJ*, 710, 1672
- Milone A. P. et al., 2014, *ApJ*, 785, 21
- Miyazaki S. et al., 2002, *PASJ*, 54, 833
- Monelli M. et al., 2013, *MNRAS*, 431, 2126
- Nelan J. E., Smith R. J., Hudson M. J., Wegner G. A., Lucey J. R., Moore S. A. W., Quinney S. J., Suntzeff N. B., 2005, *ApJ*, 632, 137
- Newman J. A. et al., 2013, *ApJS*, 208, 5
- Norris J. E., Da Costa G. S., 1995, *ApJ*, 447, 680
- Norris M. A. et al., 2008, *MNRAS*, 385, 40
- Peng E. W. et al., 2006, *ApJ*, 639, 95
- Pierce M. et al., 2006, *MNRAS*, 368, 325
- Pota V. et al., 2013a, *MNRAS*, 428, 389
- Pota V., Graham A. W., Forbes D. A., Romanowsky A. J., Brodie J. P., Strader J., 2013b, *MNRAS*, 433, 235
- Puzia T. H., Kissler-Patig M., Thomas D., Maraston C., Saglia R. P., Bender R., Goudfrooij P., Hempel M., 2005, *A&A*, 439, 997
- Rey S.-C. et al., 2007, *ApJS*, 173, 643
- Roediger J. C., Courteau S., Graves G., Schiavon R. P., 2014, *ApJS*, 210, 10
- Romanowsky A. J., Strader J., Spitler L. R., Johnson R., Brodie J. P., Forbes D. A., Ponman T., 2009, *AJ*, 137, 4956
- Salpeter E. E., 1955, *ApJ*, 121, 161
- Sánchez-Blázquez P. et al., 2006, *MNRAS*, 371, 703
- Schiavon R. P., Caldwell N., Conroy C., Graves G. J., Strader J., MacArthur L. A., Courteau S., Harding P., 2013, *ApJ*, 776, L7
- Sharina M. E., Afanasiev V. L., Puzia T. H., 2006, *MNRAS*, 372, 1259
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Smith R. J., Lucey J. R., Price J., Hudson M. J., Phillipps S., 2012, *MNRAS*, 419, 3167
- Sohn S. T., O’Connell R. W., Kundu A., Landsman W. B., Burstein D., Bohlin R. C., Frogel J. A., Rose J. A., 2006, *AJ*, 131, 866
- Spitler L. R., Larsen S. S., Strader J., Brodie J. P., Forbes D. A., Beasley M. A., 2006, *AJ*, 132, 1593
- Strader J., Smith G. H., 2008, *AJ*, 136, 1828
- Strader J., Brodie J. P., Cenarro A. J., Beasley M. A., Forbes D. A., 2005, *AJ*, 130, 1315
- Strader J., Brodie J. P., Spitler L., Beasley M. A., 2006, *AJ*, 132, 2333
- Strader J., Caldwell N., Seth A. C., 2011a, *AJ*, 142, 8
- Strader J. et al., 2011b, *ApJS*, 197, 33
- Strader J. et al., 2012, *ApJ*, 760, 87
- Strader J. et al., 2013, *ApJ*, 775, L6

- Taylor M. B., 2005, in Shopbell P., Britton M., Ebert R., eds, ASP Conf. Ser. Vol. 347, *Astronomical Data Analysis Software and Systems XIV*. Astron. Soc. Pac., San Francisco, p. 29
- Terlevich A. I., Forbes D. A., 2002, *MNRAS*, 330, 547
- Thomas D., Maraston C., Bender R., 2003, *MNRAS*, 339, 897
- Thomas D., Maraston C., Bender R., Mendes de Oliveira C., 2005, *ApJ*, 621, 673
- Tony J. L., Dressler A., Blakeslee J. P., Ajhar E. A., Fletcher A. B., Luppino G. A., Metzger M. R., Moore C. B., 2001, *ApJ*, 546, 681
- Usher C. et al., 2012, *MNRAS*, 426, 1475
- Usher C., Forbes D. A., Spitler L. R., Brodie J. P., Romanowsky A. J., Strader J., Woodley K. A., 2013, *MNRAS*, 436, 1172
- Valdes F., Gupta R., Rose J. A., Singh H. P., Bell D. J., 2004, *ApJS*, 152, 251
- VandenBerg D. A., Brogaard K., Leaman R., Casagrande L., 2013, *ApJ*, 775, 134
- Vanderbeke J. et al., 2014a, *MNRAS*, 437, 1725
- Vanderbeke J. et al., 2014b, *MNRAS*, 437, 1734
- Vazdekis A., Cenarro A. J., Gorgas J., Cardiel N., Peletier R. F., 2003, *MNRAS*, 340, 1317
- Vazdekis A., Sánchez-Blázquez P., Falcón-Barroso J., Cenarro A. J., Beasley M. A., Cardiel N., Gorgas J., Peletier R. F., 2010, *MNRAS*, 404, 1639
- Vazdekis A., Ricciardelli E., Cenarro A. J., Rivero-González J. G., Díaz-García L. A., Falcón-Barroso J., 2012, *MNRAS*, 424, 157
- Villegas D. et al., 2010, *ApJ*, 717, 603
- Woodley K. A., Gómez M., Harris W. E., Geisler D., Harris G. L. H., 2010a, *AJ*, 139, 1871
- Woodley K. A., Harris W. E., Puzia T. H., Gómez M., Harris G. L. H., Geisler D., 2010b, *ApJ*, 708, 1335
- Worthey G., Faber S. M., Gonzalez J. J., Burstein D., 1994, *ApJS*, 94, 687

## APPENDIX A: COLOUR AND SPECTRAL INDEX MEASUREMENTS

**Table A1.** Colour and spectral index measurements of spectra stacked by colour

Galaxy	$(g - i)$ (mag)	$M_i$ (mag)	CaT (Å)	Fe86 (Å)	Mg88 (Å)	Na82 (Å)	TiO89	H $\alpha$ (Å)	PaT (Å)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
NGC 1407	0.805 <sup>+0.005</sup> <sub>-0.012</sub>	-10.865 <sup>+0.363</sup> <sub>-0.012</sub>	4.672 <sup>+0.242</sup> <sub>-0.113</sub>	1.669 <sup>+0.183</sup> <sub>-0.245</sub>	0.176 <sup>+0.080</sup> <sub>-0.102</sub>	0.248 <sup>+0.191</sup> <sub>-0.241</sub>	0.9883 <sup>+0.0118</sup> <sub>-0.0064</sub>	2.270 <sup>+0.143</sup> <sub>-0.165</sub>	1.100 <sup>+0.702</sup> <sub>-0.232</sub>
NGC 1407	0.856 <sup>+0.009</sup> <sub>-0.014</sub>	-10.899 <sup>+0.185</sup> <sub>-0.216</sub>	5.455 <sup>+0.214</sup> <sub>-0.307</sub>	1.841 <sup>+0.273</sup> <sub>-0.190</sub>	0.520 <sup>+0.052</sup> <sub>-0.179</sub>	-0.165 <sup>+0.208</sup> <sub>-0.115</sub>	0.9979 <sup>+0.0070</sup> <sub>-0.0164</sub>	2.228 <sup>+0.124</sup> <sub>-0.222</sub>	1.370 <sup>+0.158</sup> <sub>-0.566</sub>
NGC 1407	0.895 <sup>+0.009</sup> <sub>-0.009</sub>	-11.153 <sup>+0.084</sup> <sub>-0.134</sub>	5.804 <sup>+0.204</sup> <sub>-0.100</sub>	2.543 <sup>+0.123</sup> <sub>-0.262</sub>	0.457 <sup>+0.090</sup> <sub>-0.050</sub>	-0.012 <sup>+0.135</sup> <sub>-0.090</sub>	0.9892 <sup>+0.0114</sup> <sub>-0.0045</sub>	1.931 <sup>+0.243</sup> <sub>-0.130</sub>	1.008 <sup>+0.239</sup> <sub>-0.237</sub>
...	...	...	...	...	...	...	...	...	...

*Notes.* The full version of this table is provided in a machine readable form in the online Supporting Information. Column (1): Galaxy. Column (2):  $(g - i)$  colour. Column (3):  $i$ -band absolute magnitude. Column (4): CaT index strength. Column (5): Fe86 index strength. Column (6): Mg88 index strength. Column (7): Na82 index strength. Column (8): TiO89 index strength. Column (9): H $\alpha$  index strength. Column (10): PaT index strength.

**Table A2.** Colour and spectral index measurements of spectra stacked by colour.

Galaxy	$(g - i)$ (mag)	$M_i$ (mag)	CaT (Å)	Fe86 (Å)	Mg88 (Å)	Na82 (Å)	TiO89	H $\alpha$ (Å)	PaT (Å)	$\sigma$ (km s <sup>-1</sup> )
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
Faint	0.8207 <sup>+0.0097</sup> <sub>-0.0089</sub>	-10.494 <sup>+0.099</sup> <sub>-0.128</sub>	5.995 <sup>+0.133</sup> <sub>-0.168</sub>	2.328 <sup>+0.122</sup> <sub>-0.138</sub>	0.436 <sup>+0.046</sup> <sub>-0.068</sub>	0.443 <sup>+0.058</sup> <sub>-0.074</sub>	0.9920 <sup>+0.0024</sup> <sub>-0.0021</sub>	1.931 <sup>+0.069</sup> <sub>-0.047</sub>	1.036 <sup>+0.133</sup> <sub>-0.204</sub>	16.82 <sup>+1.95</sup> <sub>-1.75</sub>
Faint	0.8060 <sup>+0.0060</sup> <sub>-0.0055</sub>	-9.708 <sup>+0.108</sup> <sub>-0.111</sub>	5.604 <sup>+0.155</sup> <sub>-0.177</sub>	2.171 <sup>+0.118</sup> <sub>-0.123</sub>	0.391 <sup>+0.063</sup> <sub>-0.058</sub>	0.459 <sup>+0.073</sup> <sub>-0.043</sub>	0.9916 <sup>+0.0036</sup> <sub>-0.0048</sub>	2.050 <sup>+0.050</sup> <sub>-0.138</sub>	1.116 <sup>+0.133</sup> <sub>-0.169</sub>	15.81 <sup>+0.89</sup> <sub>-1.72</sub>
Faint	0.8019 <sup>+0.0085</sup> <sub>-0.0092</sub>	-9.039 <sup>+0.104</sup> <sub>-0.098</sub>	5.560 <sup>+0.206</sup> <sub>-0.201</sub>	2.426 <sup>+0.129</sup> <sub>-0.216</sub>	0.278 <sup>+0.099</sup> <sub>-0.097</sub>	0.541 <sup>+0.058</sup> <sub>-0.140</sub>	0.9930 <sup>+0.0076</sup> <sub>-0.0081</sub>	2.104 <sup>+0.103</sup> <sub>-0.172</sub>	0.998 <sup>+0.237</sup> <sub>-0.294</sub>	13.14 <sup>+2.55</sup> <sub>-2.10</sub>
...	...	...	...	...	...	...	...	...	...	...

*Notes.* The full version of this table is provided in a machine readable form in the online Supporting Information. Column (1): Galaxy. Column (2):  $(g - i)$  colour. Column (3):  $i$ -band absolute magnitude. Column (4): CaT index strength. Column (5): Fe86 index strength. Column (6): Mg88 index strength. Column (7): Na82 index strength. Column (8): TiO89 index strength. Column (9): H $\alpha$  index strength. Column (10): PaT index strength. Column (11): Velocity dispersion.

## APPENDIX B: COLOUR AND SPECTRAL INDEX RELATIONS WITH MAGNITUDE

**Table B1.** Colour and spectral index relations with magnitude.

Sample (1)	Feature (2)	Slope (3)	Intercept (4)	$p_{\text{slope}}$ (5)	$M_i = -10$ (6)	$p_{\text{diff}}$ (7)
Faint Blue	$(g - i)$	<b><math>-0.0148^{+0.0051}_{-0.0062}</math></b>	<b><math>0.665^{+0.050}_{-0.058}</math></b>	<b>0.999</b>	$0.8128^{+0.0056}_{-0.0045}$	0.823
Bright Blue	$(g - i)$	<b><math>-0.0115^{+0.0064}_{-0.0086}</math></b>	<b><math>0.692^{+0.066}_{-0.090}</math></b>	<b>0.963</b>	$0.8074^{+0.0039}_{-0.0044}$	0.823
Faint Red	$(g - i)$	$0.0024^{+0.0051}_{-0.0055}$	$1.093^{+0.049}_{-0.052}$	0.320	$1.0688^{+0.0054}_{-0.0055}$	0.064
Bright Red	$(g - i)$	$0.0051^{+0.0040}_{-0.0033}$	$1.130^{+0.043}_{-0.036}$	0.075	$1.0789^{+0.0048}_{-0.0038}$	0.064
Faint Blue	CaT	<b><math>-0.242^{+0.137}_{-0.084}</math></b>	<b><math>3.38^{+1.34}_{-0.80}</math></b>	<b>0.959</b>	<b><math>5.797^{+0.094}_{-0.096}</math></b>	<b>1.000</b>
Bright Blue	CaT	$-0.282^{+0.174}_{-0.140}$	<b><math>2.23^{+1.84}_{-1.41}</math></b>	<b>0.955</b>	<b><math>5.046^{+0.127}_{-0.061}</math></b>	<b>1.000</b>
Faint Red	CaT	$-0.117^{+0.097}_{-0.118}$	$6.47^{+1.00}_{-1.11}$	0.876	<b><math>7.643^{+0.108}_{-0.057}</math></b>	<b>0.998</b>
Bright Red	CaT	$0.115^{+0.083}_{-0.083}$	$8.37^{+0.94}_{-0.88}$	0.070	<b><math>7.224^{+0.123}_{-0.089}</math></b>	<b>0.998</b>
Faint Blue	Fe86	$-0.105^{+0.117}_{-0.098}$	$1.24^{+1.13}_{-0.95}$	0.802	<b><math>2.291^{+0.070}_{-0.089}</math></b>	<b>0.988</b>
Bright Blue	Fe86	$-0.101^{+0.144}_{-0.136}$	$0.98^{+1.52}_{-1.44}$	0.763	<b><math>1.985^{+0.108}_{-0.118}</math></b>	<b>0.988</b>
Faint Red	Fe86	$0.081^{+0.065}_{-0.118}$	$4.54^{+0.66}_{-1.20}$	0.276	<b><math>3.729^{+0.055}_{-0.072}</math></b>	<b>0.995</b>
Bright Red	Fe86	$-0.027^{+0.097}_{-0.066}$	$3.12^{+1.09}_{-0.73}$	0.570	<b><math>3.395^{+0.127}_{-0.096}</math></b>	<b>0.995</b>
Faint Blue	Mg88	$-0.062^{+0.061}_{-0.051}$	$-0.22^{+0.60}_{-0.49}$	0.851	<b><math>0.396^{+0.041}_{-0.039}</math></b>	<b>0.997</b>
Bright Blue	Mg88	<b><math>-0.100^{+0.051}_{-0.063}</math></b>	<b><math>-0.75^{+0.54}_{-0.67}</math></b>	<b>0.963</b>	<b><math>0.250^{+0.045}_{-0.051}</math></b>	<b>0.997</b>
Faint Red	Mg88	$-0.017^{+0.043}_{-0.036}$	$0.57^{+0.43}_{-0.36}$	0.661	$0.737^{+0.036}_{-0.026}$	0.946
Bright Red	Mg88	$-0.003^{+0.028}_{-0.029}$	$0.63^{+0.32}_{-0.33}$	0.558	$0.660^{+0.042}_{-0.043}$	0.946
Faint Blue	Na82	$0.018^{+0.077}_{-0.051}$	$0.65^{+0.76}_{-0.50}$	0.315	<b><math>0.464^{+0.038}_{-0.043}</math></b>	<b>0.996</b>
Bright Blue	Na82	$-0.067^{+0.061}_{-0.065}$	$-0.41^{+0.66}_{-0.68}$	0.870	<b><math>0.263^{+0.055}_{-0.053}</math></b>	<b>0.996</b>
Faint Red	Na82	$-0.040^{+0.047}_{-0.040}$	$0.29^{+0.46}_{-0.40}$	0.801	$0.686^{+0.037}_{-0.048}$	0.432
Bright Red	Na82	$0.040^{+0.043}_{-0.053}$	$1.10^{+0.48}_{-0.60}$	0.228	$0.705^{+0.056}_{-0.069}$	0.432
Faint Blue	TiO89	$-0.0040^{+0.0033}_{-0.0035}$	$0.952^{+0.032}_{-0.035}$	0.902	$0.9920^{+0.0017}_{-0.0025}$	0.101
Bright Blue	TiO89	$0.0065^{+0.0048}_{-0.0061}$	$1.065^{+0.051}_{-0.066}$	0.152	$0.9995^{+0.0042}_{-0.0054}$	0.101
Faint Red	TiO89	$0.0000^{+0.0034}_{-0.0027}$	$1.010^{+0.034}_{-0.028}$	0.427	$1.0094^{+0.0023}_{-0.0028}$	0.939
Bright Red	TiO89	$-0.0029^{+0.0026}_{-0.0028}$	$0.973^{+0.030}_{-0.030}$	0.874	$1.0021^{+0.0037}_{-0.0037}$	0.939
Faint Blue	H $\alpha$	$0.072^{+0.124}_{-0.067}$	$2.72^{+1.23}_{-0.66}$	0.138	$1.997^{+0.027}_{-0.064}$	0.162
Bright Blue	H $\alpha$	$0.119^{+0.084}_{-0.131}$	$3.29^{+0.90}_{-1.41}$	0.176	$2.096^{+0.070}_{-0.116}$	0.162
Faint Red	H $\alpha$	$-0.106^{+0.078}_{-0.115}$	$0.91^{+0.77}_{-1.17}$	0.920	$1.970^{+0.075}_{-0.064}$	0.916
Bright Red	H $\alpha$	$-0.105^{+0.098}_{-0.077}$	$0.72^{+1.11}_{-0.88}$	0.849	$1.768^{+0.139}_{-0.105}$	0.916
Faint Blue	PaT	$-0.32^{+0.23}_{-0.10}$	$-2.2^{+2.2}_{-1.0}$	0.942	$1.053^{+0.088}_{-0.136}$	0.589
Bright Blue	PaT	$-0.06^{+0.24}_{-0.21}$	$0.4^{+2.5}_{-2.3}$	0.602	$0.976^{+0.188}_{-0.192}$	0.589
Faint Red	PaT	$-0.09^{+0.22}_{-0.13}$	$-0.2^{+2.2}_{-1.3}$	0.606	$0.677^{+0.127}_{-0.116}$	0.350
Bright Red	PaT	$-0.13^{+0.17}_{-0.12}$	$-0.5^{+1.9}_{-1.4}$	0.799	$0.764^{+0.238}_{-0.174}$	0.350

Notes. Column (1): Sample. Blue is spectra in the colour range  $0.72 < (g - i) < 0.88$ ; Red is spectra in the colour range  $1.00 < (g - i) < 1.16$ ; Faint is the spectra from NGC 3115, NGC 3377, NGC 4473 and NGC 4494; Bright is the spectra from NGC 1407, NGC 4278, NGC 4365 and NGC 4649. Column (2): spectral feature or colour. Column (3): slope of the fitted linear relation between the spectral feature in Column (2) and magnitude for the sample in Column (1). The uncertainties for Columns (3), (4) and (6) were estimated using bootstrapping. Column (4): intercept for the fitted linear relation. Column (5): probability that the slope is less than zero from bootstrapping. Column (6): spectral feature value or colour at  $M_i = -10$ . Column (7): probability that the value at  $M_i = -10$  is greater in the faint group than in bright galaxy group. Samples and features with significant slopes or differences ( $p < 0.05$  or  $p > 0.95$ ) are in bold.

## SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

**Table A1.** Colour and spectral index measurements of spectra stacked by colour.

**Table A2.** Colour and spectral index measurements of spectra stacked by colour.

(<http://mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stu2050/-/DC1>).

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the paper.

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.